Closed-form expressions for particle relative velocities induced by turbulence
(Research Note)

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ABSTRACT

In this note we present complete, closed-form expressions for random relative velocities between colliding particles of arbitrary size in nebula turbulence. These results are exact for very small particles (those with stopping times much shorter than the large eddy overturn time) and are also surprisingly accurate in complete generality (that is, also apply for particles with stopping times comparable to, or much longer than, the large eddy overturn time). We note that some previous studies may have adopted previous simple expressions, which we find to be in error regarding the size dependence in the large particle regime.

Key words. Turbulence – Dust – Planetary systems: protoplanetary disks

1. Introduction and outline

Gas in astrophysical environments is often in a turbulent state of motion, constantly affected by temporally and spatially varying accelerations from eddies having a variety of scales. A particle, due to its inertia, does not instantaneously follow the gas motions but requires a certain time in order to align with the gas motion. The particle’s interaction with the gas is captured in the definition of the stopping time of the particle (sometimes also referred to as friction time),

\[ \tau_s = \frac{3}{4c_g \rho_g} \frac{m}{\sigma} \]

where \( c_g \) and \( \rho_g \) are, respectively, the sound speed and the volume mass density of the gas, and \( m \) and \( \sigma \) the mass and projected surface area of the particle. Due to this inertial lag, a particle develops a relative velocity with respect to the gas. In addition, these lags also cause particles to acquire relative velocities among themselves.

While the general problem of calculating these relative velocities has received considerable attention in the basic fluid dynamics community (see Cuzzi & Hogan 2003 for references; henceforth CH03), the formalism most frequently used in the astrophysics community was developed by Völk et al. (1980) and Markiewicz et al. (1991) (henceforth MMV). In these works the final results are given in terms of integrals that were not solved analytically. Some workers have used simple fits to these numerical results in their models of dust coagulation; however, simple closed-form expressions for particle-particle relative velocities would help streamline these models (e.g. Suttner & Yorke 2001; Dullemond & Dominik 2005; Nomura & Nakagawa 2006; Ormel et al. 2007). Recently, CH03 obtained closed-form expressions from the MMV model for particle velocities in inertial space \( V_p \), for particle-gas relative velocities \( V_{pg} \), and for relative velocities between two identical particles \( V_{pp} \), but did not extend their results to the general case of two particles of different stopping times. Moreover, CH03 stressed the validity of their analytical results for particles with stopping times much shorter than the large eddy overturn time.

In this note we generalize the approach and results of CH03 to obtain closed-form expressions for relative velocities between particles of arbitrary, and unequal, size. In Sect. 2 we define important quantities and review previous work. In Sect. 3 we present two independent approaches for obtaining the desired closed-form solutions. In Sect. 4 we give our conclusions and a summary.

2. Definitions and previous work

Nebula gas turbulence is generally described as being composed of eddies having a range of spatial scales \( \ell \) and spatial frequencies \( k = 1/\ell \), with an energy spectrum \( E(k) \propto k^{-5/3} \) and total energy \( \frac{V_g^2}{2} \) per unit mass providing the normalization condition

\[
\frac{V_g^2}{2} = \int_{k_L}^{k_0} dk \ E(k),
\]

from which \( E(k) = \frac{V_g^2}{3k_L} (k/k_L)^{-5/3} \). The largest, or integral scale, eddies have spatial scale \( L = 1/k_L \), and the smallest, or Kolmogorov scale, eddies have spatial scale \( \eta = 1/k_\eta \). The form of \( E(k) \) given above is the inertial range expression most often assumed, with \( E(k) = 0 \) for \( k > k_\eta \) or \( k < k_L \). Völk et al. (1980) used a spectrum \( P(k) = 2E(k) \) and stipulated no smallest scale \( \eta \) for the turbulence, but Weidenschilling (1984) and MMV noted that a finite value for \( \eta > 0 \) had profound effects on the particle velocities, especially the relative velocities \( V_{pp} \) for small particles. Each eddy wavenumber \( k \) has a characteristic velocity

\[ \ell \propto k^{-5/3} \]

and a characteristic energy

\[ E(k) \propto \ell^{5/3} \]

which scales down to \( E(k) \propto \eta^{5/3} \) for small eddies. This scaling is consistent with the Kolmogorov turbulence theory.\[ \text{February 13, 2007} \]

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V(k) = \sqrt{2kE(k)} and overturn time \( t_o = \ell / v(k) = (kV(k))^{-1} \). Our standard definition of the particle Stokes number is \( St = t_o/\tau_o \), where \( \tau_o \) is the overturn time of the largest eddy, generally taken to be the local orbit period. The local turbulent intensity is described by its Reynolds number, \( Re \), defined as the ratio between the turbulent and the molecular kinematic viscosities, \( Re = \nu / \eta \). The values for \( \ell, v \) and \( t \) at the integral scale then follow from \( Re, \) e.g., \( \eta = Re^{-3/4}L \) and \( \tau_o = Re^{-1/2}t_o \). These expressions bring \( Re \) into the final expressions for particle velocities as a limit on certain integrals (cf. CH03 for more detail). In the notation of astrophysical “\( \alpha \)-models”, \( Re = \alpha_c^2 H_e / \nu = \alpha c^2 / \Omega \Delta \) where \( c, H_e \), and \( \nu \) are the sound speed, vertical scale height, and kinematic viscosity of the nebula gas and \( \Omega \) is the orbital frequency.

Völk et al. (1980) introduced the concept of “eddy classes”. Class I eddies vary slowly enough that a particle, upon entering a class I eddy, will forget its initial motion and align itself to the gas motions of the eddy before the eddy decays or the particle leaves the eddy. Class II eddies, on the other hand, have fluctuation times shorter than the particle’s stopping time \( t_o \), and fluctuate too rapidly to provide more than a small perturbation on the particle. The timescale on which an eddy decays is given by \( t_o \), while the eddy-crossing timescale is \( t_{cross} \approx \ell / v_{rel} = (kV_{rel}(k))^{-1} \), with \( v_{rel} \) the relative velocity between a grain and an eddy. For an eddy to be of class I both \( t_o \) and \( t_{cross} \) must be larger than the particle’s stopping time. The boundary between these classes occurs at \( k = k^* \) (or at \( t_o = t^* \)) which can be defined as (Völk et al. 1980, MMV):

\[
\frac{1}{t_o} = \frac{1}{\tau^*} + \frac{1}{t_{cross}} = \frac{1}{\tau^*} + k^* V_{rel}(k^*).
\]

It is important to realize that \( k^* \) (or \( \tau^* \)) is a function of stopping time \( t_o \), that is, the boundary separating the two classes is different for each particle. The different treatment for the two eddy classes \( k < k^* \) and \( k > k^* \) forms the core of the derivation of the turbulence-induced particle velocities.

All turbulent velocities in this note are statistical, root-mean-square, averaged quantities. The average inertial space particle velocity \( V_p \) is given by Eq. (6) of MMV:

\[
V_p^2 = \int_{k_h}^{\infty} dk 2E(k) \left( 1 - K^2 \right) + \int_{k_h}^{k_{crit}} dk 2E(k) \left( 1 - K \right) \left( g(\chi) + Kh(\chi) \right),
\]

in which \( K = t_o/(t_o + \tau_o) \). The \( K^2 \) term in the first integral results from the more recently preferred “\( n = 1 \)” gas velocity autocorrelation function (MMV and CH03). The functions \( g(\chi) = \chi^{-1/2} \tan^{-1}(\chi) \) and \( h(\chi) = 1/(1 + \chi^2) \) with \( \chi = K/t_o V_{rel} \) were first obtained by Völk et al. (1980).

CH03 noted that, for very small particles with \( t_o \ll \tau_o \) or \( St \ll 1 \), the second integral becomes negligible, leaving only the first integral which is analytically solvable and for which the upper limit can be extended to \( k_h \) with negligible error. Here, to generalize the approach of CH03 to particles of arbitrary size, we approximate \( h(\chi) = g(\chi) = 1 \) for all particle sizes (see CH03 Sect. 2.2.3 for supporting logic). Numerical calculations of \( h(\chi) \) and \( g(\chi) \) validate this approximation to order unity (see Appendix A), and we gain further confidence in it from a post-priori comparison with exact numerical model results. The general expression for \( V_p^2 \) is then the same as in the \( t_o \ll \tau_o \) regime, and the same analytical result is obtained, i.e. CH03,

\[
V_p^2 = \int_{k_h}^{k_{crit}} dk 2E(k) \left( 1 - K^2 \right) = V_p^2 \left( 1 - \frac{St^2(1 - Re^{-1/2})}{(St + 1)(St + Re^{-1/2})} \right).
\]

CH03 did not give this explicit result for \( V_p \), but merely noted that it was straightforward to derive it from their Eq. (19) for \( V_{pg} \) and the general relationship \( V_{pg}^2 = V_p^2 - V_s^2 \); however we will use it explicitly here.

Comparison of the predictions of this simple expression with detailed numerical results (MMV, CH03) show that it is indeed a good approximation for arbitrary \( St \). A more accurate approximation to \( E(k) \), in which the \( g \) and \( h \) functions are approximated as power-laws in \( k^* / k \), is outlined in Appendix A. Unless \( St \ll 1 \), we can neglect the Reynolds number term in Eq. (6) and obtain \( V_p = V_{pg}/\sqrt{1 + \chi St} \), a well known result (Völk et al. 1980, Cuzzi et al. 1993, Schräpler & Henning 2004) which describes the diffusivity of large particles in turbulence.

### 3. Results

#### 3.1. k-space approach

MMV (their Eq. (7)) expressed the relative velocities \( V_{pip2} \) between particles of different stopping times \( t_1 \) and \( t_2 \) as

\[
V_{pip2}^2 = V_{p1}^2 + V_{p2}^2 - 2V_{p1}V_{p2} \equiv \Delta V_{12}^2,
\]

Having already derived \( V_{p1}^2 \) (\( i = 1, 2 \)) above, we can determine \( \Delta V_{12} \) by evaluating the cross term \( V_{p1}V_{p2} \); this paper presents analytical solutions of this problem obtained in two separate ways. In this subsection we retain the wavenumber dependence; in the next subsection we transform to time variables. In Eq. (8) of MMV the cross term is given as a sum over the two particle sizes involved, which we separate here, writing \( \Delta V_{12}^2 = V_{p1}^2 + V_{p2}^2 - (V_{c1}^2 + V_{c2}^2) \), where

\[
V_{ci}^2 = \frac{2t_i}{t_1 + t_2} \left[ \int_{k_h}^{\min(k_i', k_h')} E(k)dk - \int_{k_h}^{\min(k_i', k_h')} E(k) \left( \frac{1}{1 + k_h/k_i} \right)^2 dk \right].
\]

Changing variable to \( x = k/k_h \), substituting for \( E(k) \), and converting stopping time \( t_i \) to Stokes number \( St_i = t_i/\tau_i \):

\[
V_{ci}^2 = \frac{2V_{p1}^2 t_i}{3(t_1 + t_2)} \left[ \int_{1}^{x_{i1}'} x^{-5/3}dx - \int_{1}^{x_{i2}'} \frac{St_i^2dx}{x^{5/3}(St_i + x^{-2/3})^2} \right].
\]

where we have taken, without loss of generality, \( k_{i1}' \leq k_{i2}' \). The first integral is trivial and the second integral can be solved exactly as in Eqs. (17–19) of CH03. In evaluating the specific value of the integrals above, we need a closed form for the upper limit \( x_{i1}' = k_{i1}'/k_h \). A simple prescription is readily found by inspection of Fig. 3 of CH03: \( x_{i2}' = k_{i2}'/k_h = 0.5St_i^{-3/2} + 1 \). That is, the boundary eddy for particles with stopping time \( t_i \) is that for which \( k_{i2}' \sim t_i \) until \( t_i > t_h \), beyond which it remains constant.
This is merely a convenient mathematical shorthand to keep everything in closed form. Then, repeating the analytical solution of CH03 (Eqs. (17–19)) we obtain

$$\begin{align*}
V_{ci}^2 &= \frac{V_g^2}{\frac{t_i}{(t_i + t_L)}} \left( 1 - x_1^{-2/3} \right) + \left( \frac{St_i}{1 + St_i} - \frac{St_i}{1 + St_i x_1^{2/3}} \right), \\
&= \frac{V_g^2}{\frac{t_i}{(t_i + t_L)}} \left( \left[ \frac{t_i}{t_1 + t_L} \right]^{t_i} + \left[ \frac{t_i}{t_1 + t_L} \right]^{t_L} + \left( 1 \leftrightarrow 2 \right) \right)
\end{align*}$$

(10)

This solution for the cross term is easily combined with Eq. (6) to obtain expressions for particle-particle relative velocities $\Delta V_{12}^2$. Further manipulation of these expressions may be possible, but the important point here is that $\Delta V_{12}^2$ can be expressed in closed form as function of $St_i$, $St_2$, $V_g$, and $Re$. With a few minutes of algebra, simpler expressions can be found in the limiting regimes of interest ($St_1 \ll 1, \gg 1, \ldots$) which agree well with those which we present in the next section, for analytical solutions obtained in the time domain instead of the wavenumber domain, and where an analytical solution for the boundary $k' (t^*)$ is used rather than the form for $x_1^*$ adopted above. It should be recalled that, for very small particles $t_s < t_\eta$, $x_1^*$ has an upper limit of $k_\eta/k_L = Re^{1/4}$ (see, e.g. CH03 Fig. 3).

### 3.2. t-space approach

The integrals expressing $V_{pi}^2$ and $V_{ci}^2$ are transformed into a simpler form by changing variables from $k$ to $t_i$. Since $t_i = 1/kV(k) = \left( \frac{k}{\sqrt{2\pi} A(k)} \right)^{-1}$ and $E(k) = A(k)^{-5/3}$ for a Kolmogorov power spectrum (where $A(k)$ is the normalization factor), we obtain that $E(k)dk = \frac{1}{2\pi}k^{1/2}A^{1/2}dt_i$. Now, $A = \frac{1}{2\pi}k_L^{1/2}$ from the normalization of the turbulent spectrum (Eq. (2)), $k_L = (V_L t_i)^{-1}$ with $V_L$ the velocity of the largest eddy, and $V_{ci}^2 = \frac{1}{2}V_g^2$ also by normalizing the power spectrum (see CH03). We then end up with

$$E(k)dk = \frac{1}{2t_i}V_g^2 dt_i,$$

(11)

which can be substituted into all the integrals, putting them into a simpler form. For instance, Eq. (5) becomes for particle $i$

$$V_{pi}^2 = \frac{V_g^2}{t_i} \int_{t_i}^{\infty} dt_i \left( 1 - \left( \frac{t_i}{t_i + k_L} \right)^2 \right) = \frac{V_g^2}{t_i} \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i},$$

(12)

Similarly, the cross term becomes

$$V_{ci}^2 = \frac{V_g^2}{t_i} \int_{t_i}^{\infty} dt_i \left( 1 - \left( \frac{t_i}{t_i + k_L} \right)^2 \right) = \frac{V_g^2}{t_i} \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i + t_L},$$

(13)

$$\Delta V_{12}^2 = \frac{V_g^2}{t_i} \left[ \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i} + \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i + t_L} \right] - \left( \frac{2t_i}{t_i + t_L} \right) \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i + t_L} + (1 \leftrightarrow 2),$$

(14)

With $t_{12}' = \max(t_{12}', t_{12})$ and $t_\eta \leq t_{12}' \leq t_i$ since $t^*$ refers to an eddy’s turn-over time. We now solve for $\Delta V_{12}^2$ by splitting the integral in Eq. (12) at $t_{12}'$ and subtracting the corresponding $V_{ci}$ terms from Eq. (14) to get

$$\Delta V_{12}^2 = \frac{V_g^2}{t_i} \left[ \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i} + \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i + t_L} \right] - 2 \frac{t_i}{t_i + t_L} \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i + t_L} + (1 \leftrightarrow 2),$$

(15)

where the $(1 \leftrightarrow 2)$ symbol indicates interchange between particles 1 and 2. With further manipulation and cancellation of terms, the previous expression simplifies slightly to

$$\begin{align*}
\Delta V_{12}^2 &= \frac{V_g^2}{t_i} \left[ \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i} + \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i + t_L} \right] + (1 \leftrightarrow 2) \\
&= \Delta V_{i2}^2 + \Delta V_{12}^2
\end{align*}$$

(16)

This is perhaps the most concise way to write the expressions for $\Delta V_{12}^2$. The first term we call $\Delta V_{ii}^2$ since this term involves class II (fast) eddies. If $t_{12}' = t_i$ (heavy particles) all eddies are fast and only this term remains. Conversely, if $t_{12}' = t_\eta$ (small particles) the contribution from $\Delta V_{ii}^2$ vanishes and the second term, $\Delta V_{12}^2$, determines relative velocities in the intermediate regime, $t_\eta < t_{12}' < t_i$, both terms contribute. Written in terms of the Stokes numbers these terms becomes

$$\begin{align*}
\Delta V_{i2}^2 &= \frac{V_g^2}{t_i} \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i} + (1 \leftrightarrow 2) \\
&= \frac{V_g^2}{t_i} \left[ \frac{t_i}{St_i} - \frac{St_i}{St_i + St_2} - St_i^2 - 1 \right] + (1 \leftrightarrow 2)
\end{align*}$$

(17)

$$\begin{align*}
\Delta V_{ii}^2 &= \frac{V_g^2}{t_i} \left[ t_i + \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i} + (1 \leftrightarrow 2) \\
&= \frac{V_g^2}{t_i} \left( St_i^2 - Re^{1/2} \right) + \frac{St_i^2}{St_i + St_2} - St_i^2 - 1 \right) + (1 \leftrightarrow 2)
\end{align*}$$

(18)

Note again that since $t_\eta \leq t_{12}' \leq t_i$ we also have that $Re^{1/2} \leq St_i^2 \leq 1$. Below, we will first solve for $St_i^2$, and then consider solutions for $\Delta V_{12}^2$ in various limiting cases of the particle stopping times.

#### 3.2.1. Solving for $t^*$

The relative velocity between a particle with stopping time $t_i$ and an eddy $k$, is given by Volk et al. (1980), Eq. (15):

$$V_{\text{rel}}(k)^2 = V_o^2 + 2 \int_k^k dk' E(k') \left( \frac{t_i}{t_i + k_L} \right)^2.$$  

(19)

$V_o$ is any systematic velocity component not driven by turbulence – such as due to pressure-gradient driven azimuthal headwind, the ensuing radial drift, or vertical settling under solar gravity. We can integrate this equation in the same fashion as Eq. (14) and arrive at

$$V_{\text{rel}}(k)^2 = V_o^2 + \frac{V_g^2}{t_i} \left[ \frac{t_i^2}{t_i + k_L} \right]^{t_i}_{t_i} - \frac{1}{1 + y^2} - \frac{1}{1 + y t_L},$$

(20)

in which $y = t_\eta/t_i$. Also, using the definition for $t_L$ (see text above Eq. (14)), $k^*$ can be expressed as $(k^*)^2 = (2A)_{-5/3} r^{-3} = \frac{1}{2}V_g^2 t_i r^{-3}$. Inserting the expressions for $k^*$ and $V_{\text{rel}}^2$ into Eq. (3), assuming that $V_o = 0$ for simplicity (see however Sect. 5.5), we obtain:

$$\frac{1}{k^*} \left( \frac{1}{t_\eta} - \frac{1}{t_i} \right) = V_{\text{rel}} \Rightarrow$$

(21a)
where we have defined \( y' = t'/t_b \) and \( y_l = t_b/t_e = \text{St}^{-1} \). The LHS of Eq. (21d) is negligible and the numerical solution for \( y' \) becomes \( y' \approx y_0' = 1.6 \), or \( t' \approx 1.6 t_b \). On the other hand, when \( t_e \) nears \( t_b \), the \(-1/(1 + y')\) term causes the LHS of Eq. (21d) to drop to \(-0.5\), and \( y' \to 1 \). For \( t_e > t_b \) we always have that \( t' = t_b \); i.e., for such a particle all eddies are of class 2. In Fig. 2 we compare the exact solution (dashed line) for \( t' \) with the \( t' \approx y_0'^2 t_b = 1.6 t_b \) approximation (solid line; in both cases \( t' \leq t_b \) is simply enforced), and the empirical function \( k'/k_e = 1 + \frac{1}{9} \text{St}^{-3/2} \) (dotted line; see Sect. 5.1). The exact solution for \( \Delta V_{12} \) (Eq. 16) is given in Fig. 3 both for \( t_1 \gg t_2 \) (solid curve) and for particles of equal stopping times (dashed curve). A Reynolds number of \( Re = 10^5 \) has been adopted.

### 3.3. The role of \( V_0 \): eddy-crossing effects

Systematic velocities \( V_0 \) due to vertical settling, and pressure-gradient headwinds and drifts, will occur (eg. Nakagawa et al. 1986). Because particles drift through eddies, their transit time is affected (because \( V_{rel} \) is larger) and the boundary between class I and II eddies shifts. Cuzzi et al. (1993) include this effect, due to vertical settling, in their model of particle diffusion (their Eq. (43)). The model presented here offers a generalized way of treating this effect, which we will only sketch here.

Repeating the procedure outlined in Sect. 3.2.1 but retaining the \( V_0 \) term in \( V_{rel} \) (Eq. (20)), we end up with Eq. (21d) including a correction term

\[
\frac{2}{3} y'^2 (y'^2 - 1)^2 - \frac{1}{1 + y'} = F(y') = -\frac{\text{St}}{1 + \text{St}} + \frac{1}{\text{St} \, V_0^2 \, V_g^2}, \tag{22}
\]

where we have substituted \( \text{St} = 1/y_L \). The correction term can be roughly constrained using an estimate of the systematic drift velocity \( V_0 \sim (\text{St}/(\text{St} + 1)) \beta V_K \), where \( V_K \) is the Keplerian velocity at distance \( a \) from the Sun, \( \Omega \) is the orbit frequency, and \( \beta = (H_g/a)^2 \) is a radial pressure gradient parameter; also we take \( V_0 = a_1^{1/2} c_g \) (see, eg. Nakagawa et al. 1986 or Cuzzi & Weidenschilling 2006). Then

\[
\frac{V_0}{V_g} = \frac{\text{St}}{\text{St} + 1} \frac{\beta V_K}{a_1^{1/2} c_g} = \frac{\text{St}}{\text{St} + 1} \frac{\beta \Omega}{a_1^{1/2} H_g \Omega} = \frac{\text{St}}{\text{St} + 1} \frac{1}{a_1^{1/2} \beta_{1/2}} = \frac{\text{St}}{\text{St} + 1} \left( \frac{\beta}{\alpha} \right)^{1/2}, \tag{23}
\]

and Eq. (22) becomes,

\[
F(y') = \frac{\text{St}}{1 + \text{St}} \left( \beta \alpha \frac{1}{1 + \text{St}} - 1 \right). \tag{24}
\]

Normally \( \beta \sim 2 \times 10^{-3} \) is assumed (Nakagawa et al. 1986; Cuzzi et al. 1993), but its real value, and that of \( \alpha \), are not well known. Equation (24) shows that for a given value of \( \text{St} \), \( F(y') \) increases with increasing \( \beta/\alpha \). Consequently, \( y' = t'/t_b \) is also higher (see Fig. 1). The boundary between the class I and II eddies therefore shifts to higher values of \( t' \), that is, there are less class I eddies for high \( \beta/\alpha \) and the \( t' = t_b \) upper limit (when \( t' = t_b \)) is reached at lower Stokes numbers. Inserting
In this limit all eddies are of class I and particle, the second term on the RHS of Eq. (17) is negligible.

3.4.1. Tightly coupled particles, \( t_1 \approx t_2 \), \( t_1 \gg t_2 \)

There is an extended inertial range of eddies. Recall again that \( St^* \) solutions; similar results are easily obtained from the

Moreover, we assume that \( \beta/\alpha \) to have the largest stopping time, i.e., \( t_1 \gg t_2 \) and \( t_1^* = t_1 \). 

For example, for \( \beta/\alpha = 1 \), \( St^* \) reaches its upper limit at \( St \approx 0.67 \).

In the small particle regime (\( St \ll 1 \)), however, the exact value of \( \beta/\alpha \) is unimportant since \( F(y^*) \) is always close to zero, and the \( y^* \) approximation is justified. It is only for \( \beta/\alpha \gg St^{-1} \) that the RHS of Eq. (24) starts to become significant and \( y^* > y_0^* \). This is the weakly-turbulent or non-turbulent regime where class II eddies dominate even for small particles. In practise, however, it means that eddy crossing effects are important only if turbulence is very weak and we will not treat them further in this paper.

3.4. Limiting solutions

As intuition-building examples we obtain simple, closed-form expressions for \( \Delta V_{12}^2 \) in various limiting regimes from the \( t \)-space solutions; similar results are easily obtained from the \( k \)-space solutions (Sect. 3.1). Without loss of generality we take particle 1 to have the largest stopping time, i.e., \( t_1 \geq t_2 \) and \( t_2^* = t_2 \). Moreover, we assume that \( t_1^* \ll t_2; i.e., that \( Re_1^{1/2} \gg 1 \) and there is an extended inertial range of eddies. Recall again that \( St_1^* = Re_1^{1/2} \) for \( t_1 < t_2/y_0^* \), and that \( St^* \) will not exceed 1.

3.4.1. Tightly coupled particles, \( t_1, t_2 < t_\eta \)

In this limit all eddies are of class I and \( \Delta V_{12}^2 \rightarrow \Delta V_2^2 \). For each particle, the second term on the RHS of Eq. (17) is negligible; thus

\[
\Delta V_{12}^2 = V^2 \frac{St_1 - St_2}{St_1 + St_2 \left( \frac{St_1^2}{St_1 + Re_1^{1/2}} - \frac{St_2^2}{St_2 + Re_1^{1/2}} \right)}.
\]

In the very small particle regime \( (t_1 \ll t_\eta), St_1 \ll t_\eta/t_L = Re_1^{1/2} \) and

\[
\Delta V_{12}^2 = V^2 \frac{2St_1}{t_\eta} (St_1 - St_2)^2.
\]

Since \( V^2 = 2e^2 \eta Re^{1/2} = 2\eta V^2 t_L/t_\eta \), this expression transforms directly to

\[
\Delta V_{12}^2 = \sqrt{3/2(t_1 - t_2)} V^2 t_\eta / t_\eta,
\]

In good agreement with the heuristic, although physically motivated, expression \( \Delta V_{12} = V^2(t_1 - t_2)/t_\eta \) of Weidenschilling (1984).

3.4.2. Intermediate regime, \( t_\eta \leq t_1 \leq t_L \)

If \( t_1 \) (the stopping time of the larger particle) approaches the Kolmogorov scale, two changes occur. First, the \( St_1/(St_1 + St_2) \) term in Eq. (17) now becomes linear with \( St \), since \( St_1^* \) grows proportional to \( St_1 \) (the second term is still negligible throughout most of this regime). Relative velocities therefore increase as the square-root of stopping time. Second, class II eddies also contribute to \( \Delta V_{12}^2 \) (Eq. 19). This contribution scales also with \( St_1 \), but is significantly larger and does not disappear when \( t_1 = t_2 \).

From a physical point of view, class II eddies act as small, random kicks to the particle trajectory, while two particles captured by a class I eddy are subject to the same, systematic, change in motion. Class II eddies are therefore much more effective in generating velocity differences for similar-sized particles.

In the “fully intermediate regime”, i.e., \( t_\eta \ll t_1 \ll t_L \), we can also ignore the \( Re_1^{-1/2} \) terms in Eq. (18). In addition, the \( t^*/t_\eta = y^*_a \) approximation holds. Upon writing \( St_2 = \epsilon St_1 \), Eqs. (17) (18) become linear with \( St_1 \) and we can write \( \Delta V_{12}^2 \) as (see Appendix B)

\[
\Delta V_{12}^2 = V^2 \left[ 2y_a - (1 + \epsilon) + \frac{2}{1 + \epsilon} \left( \frac{1}{1 + y_a} + \frac{\epsilon^3}{y_a + \epsilon} \right) \right] St_1,
\]

where \( \epsilon \leq 1 \) is the ratio between the stopping times and \( y_0 = 1.6 \). For \( t_1 \gg t_2 \) we then find that \( \Delta V_{12}^2 \approx 3.0V_2^2 St_1 \), while for equal particles the numerical factor goes down to \( 2.0 \).

Written in terms of stopping times the relative velocities become: 

\[
\Delta V_{12} = [1.7 \div 2.1] V_2 \sqrt{t_1/t_\eta}.
\]

This also compares well with Weidenschilling (1984) fits for this regime (who gives prefactors of 2.1 and 3.0, respectively). Note, however, that our full expressions for \( \Delta V \) (Eqs. (16) (17) (18)) also capture the behavior near the \( t_\eta \) and \( t_L \) “turning points” (see Fig. 3).

3.4.3. Heavy particles, \( t_1 > t_L \)

If \( t_1 > t_L \), \( St_1^* = 1 \) and there is no contribution from class I eddies (Eq. 17). Also, we can neglect the \( Re_1^{-1/2} \) terms in Eq. (18) and the relative velocities simply become

\[
\Delta V_{12}^2 = \Delta V_{12}^2 = V^2 \left( \frac{1}{1 + St_1} + \frac{1}{1 + St_2} \right).
\]

This result can, of course, directly be obtained from the \( V_p \) terms (Eq. (12)) since the cross-term vanishes in this regime. For small \( St_1 \) relative velocities are still \( \sim V_p \); however, if both Stokes numbers are large, the relative velocity decreases roughly with the square root of the smallest particle stopping time. Note that the linear fit of
Weidenschilling (1984) in this regime (his Eq. (15)) is inappropriate (see, however, Völk et al. (1980); Weidenschilling (1988); Weidenschilling & Cuzzi (1993); Cuzzi et al. (1993) in which a square-root fall off is advocated). Since an explicit, closed-form solution to the Völk et al. (1980) and MMV expressions for $\Delta V_{12}$ has not previously been available, many dust coagulation models (e.g., Suttner & Yorke (2001); Dullemond & Dominik (2005); Ormel et al. (2007)) have relied on the Weidenschilling (1984) fits to calculate relative velocities. Turbulent motions and relative velocities for particles in the $t_s > t_L$ regime have therefore been underestimated in these calculations. However, concerning these works, we also think no major conclusions have been affected, since the error is introduced only for large dust particles, that is, if the system is already well evolved.

### 3.5. Contour plots

Following Völk et al. (1980) and MMV we also present our results as contour plots. Figure 3A shows, for comparison, the results of MMV, obtained by numerical evaluation of the integrals involved without an inner turbulent scale ($Re \to \infty$). The next three panels of Fig. 4 show the result using our closed-form expressions derived from Eq. (16). In panel B, the $y_a^*$ approximation has been used and, like Fig. 2 of MMV (panel A), the inner scale of the turbulence is extended to infinity so that Eqs. (28, 29) apply. Somewhat systematically higher values for $\Delta V_{12}$ when compared to MMV can be explained by the CH03 approximation for $k^*$, $k^*/k_L = 0.5St^* + 1$, and also with $Re = 10^8$. Contours are drawn twice per logarithmic decade (at $\Delta V_{12}/V_g = 3 \times 10^i$ and at $10^i$) with an additional contour at $0.8$ and $1.15$.

In panels C and D of Fig. 4, a Reynolds number of $Re = 10^8$ has been adopted. For $St < 10^{-4}$, therefore, velocities are greatly suppressed since only class I eddies remain to generate relative

---

**Fig. 4.** Contour plots of particle-particle, turbulence induced, relative velocities $\Delta V_{12}$ normalized to $V_g$. —A— Numerical results of Markiewicz et al. (1991), without inner scale ($Re \to \infty$). —B— Analogous result from our closed-form expressions with the fixed $y_a^* = 1.6$ approximation (Sect. 3.2.1). —C— Like B, but with an exact solution for $y_a^*$ and with $Re = 10^8$. —D— Using the CH03 formula for $k^*$, $k^*/k_L = 0.5St^* + 1$, and also with $Re = 10^8$. Contours are drawn twice per logarithmic decade (at $\Delta V_{12}/V_g = 3 \times 10^i$ and at $10^i$) with an additional contour at 0.8 and 1.15.
velocities and relative velocities disappear completely for equal friction times. Also, the contours are much closer spaced since in this limit the velocity $\Delta V_{12}$ is proportional to $St$ (see Eq. (26)).

4. Conclusions

We have extended and, essentially, completed the work of Cuzzi & Hogan (2003), who derived explicit, closed-form expressions for particle velocities in turbulence based on the physics originally developed by Völk et al. (1980) and Markiewicz et al. (1991). Within the framework of this physics, the only approximations used here are in Eq. (6) for the particle velocities (where a posteriori comparisons with exact numerical solutions indicate the approximation is well justified) and in Eqs. (20) et seq where the systematic velocity $V_0$ is neglected to simplify calculating the boundary between eddy classes (generalizing this step should be straightforward, however). The full analytic expression for $\Delta V_{12}$ is given by Eq. (16) (or by the sum of Eqs. (17, 18)), but more simple, explicit expressions apply in restricted regimes (provided $Re^{1/2} \gg 1$):

- Equation (27), in the very small particle limit ($t_1 \ll t_2$);
- Equation (28), in the “fully intermediate” regime, i.e., for $t_2 \ll t_1 \ll t_L$;
- Equation (29), for $t_1 \geq t_L$.

Near the $t_1 = t_2$ and $t_1 = t_L$ turning points the behavior is more complex (see Fig. 3) and for accurate analytical approximations one has to revert to the full expressions for $\Delta V$ given by Eqs. (16, 17, 18).

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References

Cuzzi, J. N. & Weidenschilling, S. J. 2006, Particle-Gas Dynamics and Primary Accretion (Meteorites and the Early Solar System II), 353–381
Weidenschilling, S. J. 1988, Formation processes and timescales for meteorite parent bodies (Meteorites and the Early Solar System), 348–371

Appendix A: A more accurate closed-form solution for $V_p^2$ and all related velocities, using power-law approximations to the functions $g$ and $h$

In Sect. 2 the very simple approximation $g(\chi) = h(\chi) = 1$ was introduced for all stokes numbers $St$ and eddy scales $k$. It proves to be quite adequate for most purposes; however, as noted in Sect. 3.5 small inaccuracies remain at the 10% level because the approximation overestimates the contributions of fast eddies to $V_p^2$ and other velocity components. Figure A.1 shows the detailed behavior of the functions $g$ and $h$ for $Re = 10^8$ for Stokes numbers of $St = 0.01, 0.1,$ and $1.0$. The inflection point for all three values of $St$ is at $k = k^*$ (recall that $k^* / k_L \approx 1 + \frac{1}{2} St^{3/2}$).

Near $k^*$, the functions are well approximated by power-laws of $-1/3$ and $-3/4$, respectively, i.e., $g(k) = (k/k^*)^{-1/3}$ and $h(k) = (k/k^*)^{-3/4}$. The success of the approximation of Sect. 2 is due to the fact that the power in the weighting function $k^E(k)$ is well approximated by $k^E(k) \approx k^E(k^*)$ (dotted line; we multiply with $k$ since we compare logarithmically) decreases rapidly with increasing $k$; thus by the time the assumption $g(k) = h(k) = 1$ becomes really bad, the relative contribution of successive terms has become small. For small $St$, the weighted contribution of eddy power has already become very small even before $k \approx k^*$ (the logic of CH03). For $St = 1$ or larger, the weighting function has dropped by nearly an order of magnitude by the time $h(k)$ (the faster-decreasing function) has dropped to 0.3 (dashed-dotted line), and this seems to account for the success of our simple assumption.

This behavior can be understood from the definition of $\chi = K_L/k V_{rel}$. For $k \gg k^*$, $K_L \approx 1$ and $V_{rel} \ll V_0$ are both constant. Then, because $k_L \propto k^{-2/3}$, $\chi$ scales as $\chi \propto k^{1/3}$ and becomes large at large $k$. Since $g(\chi) = \arctan(\chi)/\chi \propto \chi^{-1/3}$ for large $\chi$, we get that $g(k) \sim k^{-1/3}$. Similarly, $h(\chi) \sim k^{-2/3}$, which is a bit shallower than the $-3/4$ exponent observed over most regions of interest (Fig. A.1). While the $-2/3$ exponent is reached at large $k$, the $-3/4$ exponent seems more appropriate at intermediate $k$.

Yet, in our subsequent analysis, we will use the large-$k$ limit for this exponent ($-2/3$) because it simplifies the math. Thus, we approximate the $g$ and $h$ behavior as follows: unity for $k < k^*$, and power laws in $k/k^*$ with exponents of $-1/3$ and $-2/3$ for

Fig. A.1. The functions $g$ (solid lines) and $h$ (dashed lines) plotted for three different Stokes numbers at a Reynolds number of $10^8$. After $k > k^*$ the functions show power-law behavior. The power spectrum (weighted by $k$) is plotted by the dotted line.
\[ k > k^*. \] Then Eq. (4) becomes

\[
V_p^2 = \int_{k_1}^{k_1} dk 2E(k)(1 - K^2) + \int_{k_1}^{\infty} dk 2E(k)(1 - K) \left( \frac{k}{k_1} \right)^{-1/3} + K \left( \frac{k}{k_1} \right)^{-2/3} . \tag{A.1}
\]

Where we have that \( k_1 \leq k^* \leq k_1 \) such that in the case of very small or very large particles one of the integrals vanishes (Sect. 3.2.1). Since the approximation \( \varphi = h = 1 \) still holds for \( k < k^* \) (or for \( t > t_f \) the velocities resulting due to class 1 eddies (Eq. (11)) are not affected; the new approximation only affects Eq. (18). By writing \( K = St/(St + x^{2/3}) \), \( E(k) = E_1(k/k_1)^{-5/3} \propto x^{-3/3} \) with \( x = k/k_1 \) the solution to Eq. (A.1) involves integrals of the form

\[
\int dx x^{-3/3+3/p} \left( \frac{St}{St + x^{2/3}} \right)^{\eta=1,2} \tag{A.2}
\]

with \( n = 2 \) for the \( K^2 \) term and \( p = -1/3 \) or \(-2/3 \). These integrals can be solved analytically. Going to “space”, however, gives somewhat cleaner solutions and we will find on how that approach and show how it affects relative velocities, i.e., \( \Delta V_{HI} \). After the change of variables \( (t_f/t^* = (k/k_1)^{2/3}) \) the second term of Eq. (A.1) becomes

\[
\frac{V_p^2}{t_L} \int_{t_1}^{\infty} dt \left( 1 - K \right) \left( \frac{t_f}{t} \right)^{1/2} + \left( 1 - K \right) K \left( \frac{t_f}{t} \right) . \tag{A.3}
\]

We now introduce the dimensionless variable \( y = t_f/t_1 \) (cf. Eq. (21)). Then \( t_f/t^* = y/y^* \) with \( y^* = t_f/t_1 \). Also \( K = 1/(1 + y) \) and \( 1 - K = y/(1 + y) \) and Eq. (A.3) becomes

\[
\frac{V_p^2}{t_L} \int_{1/y_1}^{\infty} dy \left( y^* \right)^{-1/2} \frac{y^{3/2}}{1 + y} + \left( y^* \right)^{-1} \frac{y^2}{(1 + y)^2} \tag{A.4}
\]

\[
= \frac{V_p^2}{t_L} \left( y^* \right)^{-1/2} \left[ I_0(y) \right]_{y_1/y_0}^{y} + \left( y^* \right)^{-1} \left[ I_1(y) \right]_{y_1/y_0}^{y} \tag{A.5}
\]

in which the functions \( I_0(y) \) and \( I_1(y) \) are defined as

\[
I_0(y) \equiv \int_0^y \frac{dz}{1 + z} = \left( 2/3 \right) y - 2 \text{arctan}(\sqrt{y}) + 2 \text{arctan}(\sqrt{y}) , \tag{A.6}
\]

\[
I_1(y) \equiv \int_0^y \frac{z^2}{(1 + z)^2} = \frac{(2 + y) y}{1 + y} - 2 \text{log}(1 + y) + 2 \text{arctan}(\sqrt{y}) \tag{A.7}
\]

The expressions for \( \Delta V_{HI} \) now consist of several contributions. First, \( I_0(y) \) and \( I_1(y) \) are evaluated at both the upper \((y^*)\) and lower \((y_0)\) limits. This must be done for both particles 1 and 2, because the \( \Delta V_{HI} \) term (Eq. (18)) has separate contributions from each particle. For the particle of highest friction time (say this is \( t_1 \)) the power-law approximation for \( g \) and \( h \) holds over the range \( \Delta V_{HI} \) is calculated, i.e., \( t_s \leq t_1 \leq t_1^* \). However, for the second particle the power-law approximation only holds for \( t_2 \leq t_1^* \), while for the remaining range over which the integral in \( \Delta V_{HI} \) is evaluated, i.e., \( t_2^* \leq t_2 \leq t_1^* \), the \( g = h = 1 \) approximation applies.

This gives us several terms that contribute to \( \Delta V_{HI} \). Collecting these terms, the new expression for \( \Delta V_{HI} \) becomes

\[
\Delta V_{HI} = \frac{V_p^2}{t_L} \left\{ t_1 \left( \frac{t_1}{t_1} \right)^{-1/2} \left[ I_0(y) \right]_{y_0}^{y_1} + \left( 1 \leftrightarrow 2 \right) \right. \tag{B.1}
\]

\[
+ t_1 \left( \frac{t_1}{t_1} \right)^{-1} \left[ I_1(y) \right]_{y_0}^{y_1} + \left( 1 \leftrightarrow 2 \right) \right. \tag{B.2}
\]

\[
+ \left. \left\{ t_1 + \frac{t_f^2}{t_2 + t_1} t_1^* \right\}, \quad (t_1 \geq t_2). \tag{A.8}
\]

Although still fully analytical, this more accurate expression for \( \Delta V_{HI} \) is also more complicated and we did not present it in the main body of the paper. Equation (A.8) is useful, however, for readers whose applications demand this higher level of accuracy.

### Appendix B: Derivation of Eq. (28)

We consider the limiting case of \( t_f \ll t_1 \ll t_L \). The \( y^*_1 \) approximation for \( St_{12} \) then holds, i.e., \( St_{21} \approx y^*_0 St_{11} \) with \( y_0 = 1.6 \). We will now argue that we can neglect the \( Re^{-1/2} \) terms in Eq. (18). For particle 1 this is obvious since \( St_1 \gg Re^{-1/2} \). The last term (where \( Re^{-1/2} \) is in the denominator) then becomes simply \(-St_1 \). However, for the interchange term a similar approximation

\[
\frac{St_2^2}{St_2 + Re^{-1/2}} \approx St_2, \tag{B.1}
\]

is not that obvious since we have not put a constraint on \( St_2 \). For example, if \( St_2 \ll Re^{-1/2} \) the \( Re^{-1/2} \) term dominates the denominator. However, in that case this term and its approximation are small anyway compared to \(-St_1 \), such that by making the approximation in Eq. (B.1) our final result is not affected. Similarly, if \( St_2 \sim Re^{-1/2} \), Eq. (B.1) (which goes to \( \sim 1/Re^{-1/2} \)) or its approximation \( \sim Re^{-1/2} \) are insignificant since \( St_1 \gg Re^{-1/2} \). Only if \( \epsilon \sim 1 \), i.e., \( St_2 \gg Re^{-1/2} \), does the \( St_2 \) term matter, but then the approximation in Eq. (B.1) is well justified. All terms in Eq. (18) are then linear in Stokes and we can reduce it to

\[
\frac{\Delta V_{HI}}{V_p^2} = \left( 2y^*_0 - (1 - \epsilon) + \frac{1}{1 + y^*_0} \right) \frac{e^2}{\epsilon + y^*_0} \tag{B.2}
\]

with \( \epsilon = St_2/Sp_\leq 1 \). Similarly, Eq. (12) becomes

\[
\frac{\Delta V_{HI}}{V_p^2} = \frac{1 - \epsilon}{1 + \epsilon} \left( \frac{1}{1 + y^*_0} \right) \frac{e^2}{y^*_0 + \epsilon} \tag{B.3}
\]

Combining these expressions and collecting the \( 1/(1 + y^*_0) \) and \( e^2/(y^*_0 + \epsilon) \) terms then gives Eq. (28).