The HI mass function in Ursa Major

Marc Verheijen

NRAO, P.O. Box O, Socorro, NM 87801, U.S.A.

Neil Trentham

Institute of Astronomy, Madingley Road, Cambridge CB3 0EZ, U.K.

Brent Tully

Institute for Astronomy, 2680 Woodlawn Drive, Honolulu, HI 96822, U.S.A.

Martin Zwaan

Kapteyn Institute, Postbus 800, 9700 AV, Groningen, The Netherlands

Abstract. A deep blind HI survey of the nearby spiral rich Ursa Major cluster has been performed with the VLA in parallel with a wide-field R-band imaging campaign of most VLA fields with the CFHT. The goal is to measure the slope of the HI mass function (HIMF) down to HI masses of $10^7 M_{\odot}$ as well as the slope of the faint-end of the luminosity function (LF) down to $M(R)=-10$. The VLA survey sampled 16% of the Ursa Major cluster volume. Both the HIMF and the LF in Ursa Major are nearly flat at their faint-ends. No free-floating HI clouds were detected.

1. Motivation

The slope of the HI mass function (HIMF) at the low mass end is a much debated topic. It is related to the distribution of primordial density fluctuations and it helps constraining galaxy formation scenarios. In case a large undiscovered population of HI rich dwarf galaxies does exist, it may contain a non-negligible fraction of the baryon content of the universe and provide fuel for the star formation process in more massive disk galaxies through their capture, infall and accretion. Furthermore, the slope of the HIMF ties in directly with the recently launched hypothesis that the population of High Velocity Clouds (HVCs) as cataloged by Wakker & Van Woerden (1991), apparently associated with the Milky Way, might actually be of an extragalactic nature, being the primordial remnants of the formation of the Local Group (Blitz et al. 1999).

Despite its importance, the slope of the HIMF at the low-mass end has not been measured with any accuracy down to HI masses of $10^7 M_{\odot}$. In general, the statistics in the $10^8 - 10^7 M_{\odot}$ mass range are too poor, with typically 1 or 2 galaxies in the lowest HI mass bin. The Arecibo HI Strip Survey (AHISS) of Zwaan et al. (1997), still the deepest survey to date, indicates a slope in
Figure 1. The layout of the VLA pointings in the Ursa Major cluster. Note that both crowded and empty regions in the cluster are sampled. Circles: previously known cluster members (solid: complete sample, open: galaxies intrinsically fainter than the SMC). Stars: Newly discovered HI dwarfs (solid: inside cluster velocity range, open: outside velocity range). Triangles: optically selected cluster candidates (solid: almost certain member, open: probably background objects). The C-fields were pointed observations toward interesting galaxies and were omitted from the HIMF analysis. The large circles indicate the size of the FWHM of the VLA primary beam.
the range $-1.1$ to $-1.3$ with only 1 galaxy in the last bin. From the AHISS data, combined with an independent deep Arecibo Slice Survey, Schneider et al. (1998) claimed a sharp steepening of the HIMF below $10^8 M_\odot$ with only 2 galaxies in the last bin. A first HIMF comprising 263 galaxies from the HIPASS project (Kilborn et al. 1999) shows a slope of $-1.3$ with only 3 galaxies in two bins below an HI mass of $10^8 M_\odot$. At this meeting in Guanajuato, an updated slope of roughly $-1.4$ was reported from the HIPASS survey by Staveley-Smith and a slope as steep as $-1.7$ was suggested by Schneider from the new Arecibo Dual-Beam Survey (ADBS). It should be noted that the volume corrections of all these surveys are hampered by uncertain corrections for large scale structure in the Universe. These corrections should account for the fact that we are located in a local overdensity and that dwarf galaxies are more easily detected nearby than in the distant Universe, introducing a bias toward a steeper slope.

From all this one can conclude that the sensitivity of recent surveys is simply insufficient and deep observations of a well defined, significantly overdense region with field-like characteristics of its galaxy population are needed to obtain the required statistics. It was realized that the sensitivity of the VLA in its D-configuration in combination with the overdensity of the nearby spiral rich Ursa Major (UMa) cluster would provide sufficient detections per HI mass-bin down to $10^7 M_\odot$ to reliably measure the slope of the HI mass function at the low-mass end. Furthermore, all galaxies in the UMa volume would be at the same distance, eliminating the effects of large scale structure along the depth of the survey. The VLA would also be sensitive enough to detect the UMa equivalent of the proposed Local Group population of extragalactic HVCs with typical HI masses of a few times $10^7 M_\odot$, as initially suggested by Blitz et al. (1999).

2. The Ursa Major cluster

The Ursa Major cluster at a distance of 18 Mpc spans about 15 degrees on the sky and comprises some 79 Mpc$^3$. It is a gravitationally bound overdensity located in the Supergalactic plane. However, it contains no ellipticals, just a dozen lenticulars and has a velocity dispersion of only 150 km·s$^{-1}$, which corresponds to a crossing time of roughly a Hubble time (Tully et al. 1996). Its morphological mix is similar to lower density field regions. This makes the Ursa Major region an especially suitable environment to study the low-mass end of the HIMF. For instance, if the slope of the HIMF would indeed be as steep as $-1.7$ as suggested by Schneider in this meeting, there would be about 2900 HI objects in the Ursa Major volume with HI masses above $10^7 M_\odot$, based on the high-mass end of the HIMF of the 80 known cluster members.

3. The VLA blind HI survey

A VLA blind HI survey of the Ursa Major cluster and CFHT imaging of most of the VLA fields in the R-band with the wide-field camera has been performed. The optical campaign was aimed at measuring the faint end of the Luminosity Function (LF) while it would provide optical morphologies and luminosities of the HI detected dwarfs. In return, the HI detections could provide redshifts for the optically selected cluster candidates.
Figure 2. The HIMF in Ursa Major as measured with the VLA. **Upper panel:** Horizontal bars show the bin widths, vertical error-bars indicate the Poisson noise. The positions of the points along the horizontal bars indicate the average HI mass of the galaxies within each bin. The solid line shows the volume probed as a function of HI mass, assuming that the detections follow an HI mass-line width relation $\Delta V \propto M_{\text{HI}}^{1/3}$ and the profile shapes at the low mass end are Gaussian. The dots indicate the actually applied $1/V_{\text{max}}$ for individual galaxies, taking into account their measured HI masses and velocity widths. Their close adherence to the solid line indicates the validity of the assumptions that go into the analytically constructed volume corrections. **Lower panel:** Number of detected galaxies per bin.
A cross-pattern of 22×32 pointings was observed, with pointing centers separated by a primary beam FWHM as indicated in Figure 1. Both crowded and empty regions were sampled. Given the low velocity dispersion of the cluster, a 3.125 MHz bandwidth was sufficient to cover the entire velocity range with a resolution of 10.3 km s⁻¹ after Hanning smoothing. A 70 minutes integration time per pointing resulted in a rms noise of 0.79 mJy beam⁻¹ which corresponds to a minimum detectable HI mass of 5×10⁶ M⊙ at 6-sigma at the beam centers at a resolution of 45 s by 10 km s⁻¹. The data were flagged and calibrated in AIPS, mosaiced and cleaned in MIRIAD and further analyzed in GIPSY. Within the mosaiced images about 16% of the entire cluster volume was sampled. Currently, the full North-South and half of the East-West strip is processed.

4. Volume corrections

Given the depth of the cluster and the primary beam attenuation of the VLA, galaxies with small HI masses and broad profiles can not be detected throughout the entire survey volume. Therefore, a simple 1/V_{max} volume correction was applied where V_{max} is the volume in which a dwarf galaxy with a particular HI mass M_{HI} and line width W_{20} could have been detected. In an analytical approach, one can for each HI mass infer a line width according to an empirical HI mass-line width relation: 

\[ W_{20} = 0.16 \times \frac{M_{HI}^{1/3}}{M_{HI}} \]

where \( W_{20} \) is the line width at 20% of peak flux \( F_{peak} \). \( W_{20} \) relates to the FWHM \( \Delta V \) of a Gaussian profile according to 

\[ W_{20} = \Delta V \cdot \frac{\sqrt{\ln 5}}{\ln 2} \]

For a Gaussian profile, \( \int F \, dv = F_{peak} \cdot W_{20} \cdot \frac{\sqrt{\pi}}{4\ln 5} \) where \( F_{peak} = 6 \cdot \sigma_{R} \) required for a detection and \( \sigma_{R} \) is the noise at velocity resolution \( R \).

Detectability is maximized under optimal smoothing conditions (\( R = \Delta V \)) and the noise \( \sigma_{R} \) which corresponds to a situation of optimal velocity smoothing is calculated according to 

\[ \sigma_{R} = \sigma_{R=10.3} \cdot \sqrt{10.3/R} \] in which \( \sigma_{R=10.3} \) is the noise in the highest velocity resolution data. Under this formalism, a maximum distance \( D_{max} \) at which an HI mass \( M_{HI} \) can be detected is calculated according to

\[ M_{HI} = 2.36 \times 10^{5} \cdot D_{max}^{2} \cdot \sigma_{R=10.3} \cdot \sqrt{W_{20}} \cdot \sqrt{\frac{\pi}{10.3} \cdot \frac{4\ln 5}{\ln 5} \cdot \frac{(\ln 5/\ln 2)^{1/4}}} \]

Of course, \( D_{max} \) varies across the primary beam due to its attenuation characteristics which makes \( \sigma_{R} \) position dependent. Considering the near and far side of the cluster volume, \( V_{max} \) can easily be calculated from \( D_{max} \) and \( \sigma_{R} \). This analytical approach is used to draw the solid line in Figure 2. Ultimately, 1/V_{max} was calculated and applied for each individual galaxy using its measured HI mass and line width.

5. Results

The data cubes were smoothed in velocity to resolutions \( R \) of 20, 30, 40, 60 and 80 km s⁻¹ and at each velocity resolution, the noise was determined and a 6-sigma clip was applied to find the HI emission. In total 32 HI detections were made, all of them have an optical counterpart in the CFHT images. All 19 previously known cluster members within the imaged areas were detected, including all the S0 systems. Only 13 new objects were discovered in 16% of the UMa volume.
Figure 3. The Ursa Major HIMF in perspective. The volume corrected HIMF in Ursa Major as measured by the VLA is plotted as points with errorbars as explained in Figure 2. The dark-gray histogram shows the optically selected HIMF to which the VLA HIMF was normalized at the high mass end. The cross-hatched part shows contributions from galaxies intrinsically fainter than the SMC. The light-gray histogram shows the expected distribution of the UMa analogs of the proposed Local Group HVCs if they would have similar space densities in UMa as in the Local Group. The mass distribution is based on the assumption that the HVCs are extragalactic and virialized with a dark matter fraction of 95%. They seem to predict a HIMF slope of $-1.7$ in UMa (dashed line). The middle-gray histogram is a similar expected distribution of UMa analogs of the Compact HVCs identified by Braun & Burton (1999) in case these too are Local Group objects at a common distance of 650 kpc.
After applying volume corrections for the individual galaxies, the HIMF in Ursa Major as measured by the VLA is plotted in the upper panel of Figure 2. The analytical and idealized relation between $M_{\text{HI}}$ and $V_{\text{max}}$ is plotted as a solid line in the upper panel which shows that galaxies with $M_{\text{HI}}>10^{8.75}M_{\odot}$ can be detected throughout the entire imaged volume while $1/V_{\text{max}}$ becomes very large for galaxies with $M_{\text{HI}}<10^{7}M_{\odot}$.

The HI mass function in Ursa Major turns out to be nearly flat with 4 galaxies in the lowest HI mass bin. In accordance with the HIMF, the LF from the CFHT data is also nearly flat ($\alpha = -1.1$) down to $M(R) = -10$.

Figure 3 shows the VLA derived HIMF in comparison with the optically selected HIMF and the expectations for HVCs in the Ursa Major volume. The dark-gray histogram shows the previously known, optically selected UMa HIMF to which the VLA HIMF was normalized at the high mass end.

The light-gray histogram shows the expected distribution of HI masses of the population of Local Group HVCs, under the assumptions that they are indeed of an extragalactic nature and virialized with a dark matter fraction of 95%, and that Ursa Major would have the same space density of HVCs as the Local Group for which a volume of 15 Mpc$^3$ was adopted. Such a population of extragalactic HVCs would give rise to a slope of $-1.7$ as indicated by the dashed line which is highly inconsistent with the actually measured data points.

The middle-gray histogram shows the expected distribution of HI masses of the recently identified population of compact HVCs (Braun & Burton 1999) in case they were located in the Local Group at a common distance of 650 kpc, and again, if the space density in UMa would be the same as in the Local Group.

In conclusion, although this is a preliminary analysis of 85% of the data, the HI mass function in the Ursa Major cluster is flat down to HI masses of $10^7$ M$_\odot$ and a slope as steep as $-1.7$ is certainly ruled out in UMa; it would have predicted 30 times more $10^{7.5}$ M$_\odot$ objects (well above the detection limit) than observed with the VLA. No free-floating HI clouds were detected nor the UMa analogs of the hypothesized Local Group extragalactic HVCs. Their space density in UMa would have to be a factor $\sim 100$ smaller than in the Local Group.

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References

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