The X-ray virial relations for relaxed lensing clusters observed with Chandra

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ABSTRACT
We examine the relations linking mass, X-ray temperature and bolometric luminosity for a sample of luminous, relatively relaxed clusters of galaxies observed with the Chandra Observatory, for which independent confirmation of the mass results is available from gravitational lensing studies. Within radii corresponding to a fixed overdensity \( \Delta = 2500 \) with respect to the critical density at the redshifts of the clusters, the observed temperature profiles, scaled in units of \( T_{2500} \) and \( r_{2500} \), exhibit an approximately universal form which rises within \( r \sim 0.3 r_{2500} \) and then remains approximately constant up to \( r_{2500} \). We obtain best-fit slopes for the mass-temperature and temperature-luminosity relations consistent with the predictions from simple scaling arguments \( \text{i.e.} \; M_{2500} \propto T_{2500}^{2/3} \) and \( L_{2500} \propto T_{2500}^{3} \), respectively. We confirm the presence of a systematic offset of \( \sim 40 \) per cent between the normalizations of the observed and predicted mass-temperature relations for both SCDM and ΛCDM cosmologies.

Key words: X-rays: galaxies: clusters – galaxies: clusters: general – gravitational lensing – cosmological parameters

1 INTRODUCTION
The spatial distribution, mass function and redshift evolution of clusters of galaxies are sensitive functions of cosmology. The space density \( n(M, z) \) of clusters predicted by analytical models \( \text{e.g.} \; \text{Press \\& Schechter} \; 1974; \; \text{Lacey \\& Cole} \; 1993; \; \text{Sheth, Mo \\& Tormen} \; 2001) \) and numerical simulations \( \text{e.g.} \; \text{Navarro, Frenk \\& White} \; 1995; \; \text{Eke, Cole \\& Frenk} \; 1996; \; \text{Jenkins et al.} \; 2000; \; \text{Bode et al.} \; 2001) \) can be related to (more easily) observable properties such as the X-ray temperatures and luminosities of clusters via simple scaling relations. Assuming that the X-ray gas in clusters is virialized and in hydrostatic equilibrium, the mass, \( M_{\Delta} \), within radius \( r_{\Delta} \) (inside which the mean mass density is \( \Delta \) times the critical density, \( \rho_{c}(z), \) at that epoch) is related to the mean mass-weighted temperature within that radius, \( T_{\Delta} \), by \( E(z)M_{\Delta} \propto T_{\Delta}^{3/2} \). Here, \( E(z) = H(z)/H_{0} = (1 + z)\sqrt{(1 + \zeta\Omega_{m} + \Omega_{\Lambda}/(1 + z)^{2} - \Omega_{\Lambda})} \), where \( H(z) \) is the redshift-dependent Hubble Constant \( \text{e.g.} \; \text{Bryan \\& Norman} \; 1998 \). Since the X-rays from rich clusters are primarily bremsstrahlung emission, one can also show that \( L_{\Delta}/E(z) \propto T_{\Delta}^{3} \), where \( L_{\Delta} \) is the bolometric luminosity from within radius \( r_{\Delta} \). The validity of these simple scaling relations is supported by numerical simulations \( \text{e.g.} \; \text{Evrard, Metzler \\& Navarro} \; 1996; \; \text{Bryan \\& Norman} \; 1998; \; \text{Thomas et al.} \; 2001; \; \text{Mathiesen \\& Evrard} \; 2001 \), although the normalization of the mass-temperature relation exhibits some variation from study to study (the normalization of Bryan & Norman 1998 is \( 17 \) per cent higher than that of Evrard, Metzler & Navarro 1996 for \( \Delta = 250 \); see also Table 1 of Afshordi & Cen 2001). The normalization of the luminosity-temperature relation is more difficult to predict due to the potentially complex physics of the X-ray gas in the innermost regions of clusters from where the bulk of the X-ray luminosity arises.

Recent observational determinations of the mass-temperature relation, based on ASCA and ROSAT data for relatively hot \( (kT \gtrsim 3 - 4 \text{ keV}) \) clusters \( \text{e.g.} \; \text{Horner, Mushotzky \\& Scharf} \; 1999; \; \text{Nevalainen, Markevitch \\& Forman} \; 2000; \; \text{Finoguenov, Reiprich \\& Böhringer} \; 2001 \) have recovered a slope consistent with the simple scaling-law predictions, although the observed normalizations are typically \( \sim 40 \) per cent lower than predicted by the simulations of Evrard, Metzler & Navarro (1996) for a standard cold dark matter (SCDM) cosmology. For clusters at lower temperatures, some steepening of the mass-temperature relation is inferred (Nevalainen et al. 2000, Finoguenov et al. 2001). Studies of the luminosity-temperature relation \( \text{e.g.} \; \text{White, Jones \\& Forman} \; 1997; \; \text{Allen \\& Fabian} \; 1998; \; \text{Markevitch} \; 1998; \; \text{Arnaud \\& Evrard} \; 1999 \) have generally measured \( L_{\text{bol}} \propto T^{3} \), whereas theory predicts \( L_{\text{bol}} \propto T^{2} \). This has been taken as evidence for significant pre-heating and/or cooling in cluster cores \( \text{e.g.} \; \text{Kaiser} \; 1991; \; \text{Evrard \\& Henry} \; 1991; \; \text{Cavaliere, Menci \\& Tozzi} \; 1997; \; \text{Pearce et al.} \; 2000; \; \text{Bialek, Evrard \\& Mohr} \; 2001 \). Allen & Fabian (1998) have shown that for hot \( (kT \gtrsim 5 \text{ keV}) \), relaxed clusters \( L_{\text{bol}} \propto T^{2} \).
Table 1. Summary of the Chandra observations.

<table>
<thead>
<tr>
<th>Cluster</th>
<th>Redshift</th>
<th>Date</th>
<th>Net Exposure</th>
</tr>
</thead>
<tbody>
<tr>
<td>PKS0745-191</td>
<td>0.103</td>
<td>2001 Jun 16</td>
<td>17.9</td>
</tr>
<tr>
<td>Abell 2390</td>
<td>0.230</td>
<td>1999 Nov 7</td>
<td>9.1</td>
</tr>
<tr>
<td>Abell 1835</td>
<td>0.252</td>
<td>1999 Dec 12</td>
<td>19.6</td>
</tr>
<tr>
<td>MS2137-253</td>
<td>0.313</td>
<td>1999 Nov 18</td>
<td>20.6</td>
</tr>
<tr>
<td>RXJ1347-1145(1)</td>
<td>0.451</td>
<td>2000 Mar 05</td>
<td>8.9</td>
</tr>
<tr>
<td>RXJ1347-1145(2)</td>
<td>0.451</td>
<td>2000 Apr 29</td>
<td>10.0</td>
</tr>
<tr>
<td>3C295</td>
<td>0.461</td>
<td>1999 Aug 30</td>
<td>17.0</td>
</tr>
</tbody>
</table>

is recovered once the effects of cool, central components are accounted for in the spectral X-ray analysis, suggesting (in agreement with the later mass-temperature results) that pre-heating may only significantly affect the properties of cooler, less-luminous clusters.

A major goal of studies with the new generation of X-ray missions including the Chandra Observatory and XMM-Newton, which permit the first direct spatially-resolved X-ray spectroscopy of hot, distant clusters, is the verification of the NFW (Navarro, Frenk & White 1997) halo mass profile, incorporating the Fe-L calculations of Liedhal, Osterheld & Goldstein (1995), and the photoelectric absorption models of Balucinska-Church & McCammon (1999). Two separate models were applied to the data, the first of which was fitted to each annular spectrum individually in order to measure the projected temperature profiles. The second model was applied to all annuli simultaneously, in order to determine the deprojected temperature profiles under the assumption of spherical symmetry. Only data in the 0.5 – 7.0 keV range were used.

For the mass modelling, azimuthally-averaged surface brightness profiles were constructed from background subtracted, flat-fielded images with a 0.984 x 0.984 arcsec$^2$ pixel scale (2 x 2 raw detector pixels). When combined with the deprojected spectral temperature profiles, the surface brightness profiles can be used to determine the X-ray gas mass and total mass profiles in the clusters. For this analysis we have used an enhanced version of the image deprojection code described by White, Jones & Forman (1997) with distances calculated using the code of Kayser, Helbig & Schramm (1997). We have parameterized the mass profiles using a Navarro, Frenk & White (1997; hereafter NFW) model with $\rho(r) = \rho_c(z) \delta_c / [(r/r_s)^{1} + (1 + r/r_s)^{3}]^{5}$, where $\rho(r)$ is the mass density, $\rho_c(z) = 3H(z)^2/8\pi G$ is the critical density for closure at redshift $z$, and $\delta_c = 200c^3/3\ln[1 + c] - c/(1 + c)$. The normalizations of the mass profiles may also be expressed in terms of an equivalent velocity dispersion, $\sigma = \sqrt{50\rho_c H(z)}$ (with $r_s$ in units of Mpc). The best-fit NFW model parameter values and 68 per cent confidence limits are summarized in Table 2.

In determining the results on the virial properties, we adopt $\Delta = 2500$, since $r_{2500}$ is well-matched to the outermost radii at which reliable temperature measurements can be made from the Chandra S3 data. (The $r_{2500}$ values for the NFW models are determined numerically, with confidence limits calculated using the $\chi^2$ grids. Note that $r_{2500}$ varies from 0.26 – 0.33$r_{200}$ for the clusters in the present sample. We define $kT_{2500}$, the mean mass-weighted temperature within $r_{2500}$, as $kT_{2500} = \sum_{i=1}^{n} m_{gas,i} kT_i / \sum_{i=1}^{n} m_{gas,i}$, where $m_{gas,i}$ and $kT_i$ are the gas mass and temperature (in keV) in each radial shell for which an independent spec-

* For RXJ1347-1145, the data from the southeast quadrant of the cluster were excluded due to ongoing merger activity in that region; Allen et al. (2001c).

† The observed surface brightness profile and a particular parameterized mass model are together used to predict the temperature profile of the X-ray gas. (We use the median temperature profile determined from 100 Monte-Carlo simulations. The outermost pressure is fixed using an iterative technique which ensures a smooth pressure gradient in these regions.) The predicted temperature profile is rebinned to the same binning as the projected/deprojected spectra and compared with the observed spectral deprojection results. The $\chi^2$ difference between the observed and predicted temperature profiles is then calculated. The parameters for the mass model are stepped through a regular grid of values in the $r_s$-$\sigma$ plane to determine the best-fit values and 68 per cent confidence limits. Spherical symmetry and hydrostatic equilibrium are assumed throughout.

2 OBSERVATIONS AND DATA ANALYSIS

The Chandra observations were carried out using the back-illuminated S3 detector on the Advanced CCD Imaging Spectrometer (ACIS) between 1999 August 30 and 2001 June 16. For our analysis we have used the the level-2 event lists provided by the standard Chandra pipeline processing. These lists were cleaned for periods of background flaring using the CIAO software package resulting in the net exposure times summarized in Table 1.

The Chandra data have been analysed using the methods described by Allen et al. (2001b,c) and Schmidt, Allen & Fabian (2001). In brief, concentric annular spectra were extracted from the cleaned event lists, centred on the peaks of the X-ray emission from the clusters. The spectra were analysed using XSPEC (version 11.0: Arnaud 1996), the MEKAL plasma emission code (Kaastra & Mewe 1993; incorporating the Fe-L calculations of Liedhal, Osterheld & Goldstein 1995), and the photoelectric absorption models of Balucinska-Church & McCammon (1999). Two separate models were applied to the data, the first of which was fitted to each annular spectrum individually in order to measure the projected temperature profiles. The second model was applied to all annuli simultaneously, in order to determine the deprojected temperature profiles under the assumption of spherical symmetry. Only data in the 0.5 – 7.0 keV range were used.

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Table 2. The best-fit parameter values and 68 per cent ($\Delta \chi^2 = 1.0$) confidence limits for the NFW mass models. $r_s$ values are in units of Mpc and $\sigma$ values in $\text{km s}^{-1}$.

<table>
<thead>
<tr>
<th>Cluster</th>
<th>$r_s$</th>
<th>$\sigma$</th>
<th>$r_s$</th>
<th>$\sigma$</th>
</tr>
</thead>
<tbody>
<tr>
<td>PKS0745-191</td>
<td>0.85$^{+0.12}_{-0.18}$</td>
<td>3.66$^{+0.53}_{-0.26}$</td>
<td>1275$^{+75}_{-125}$</td>
<td>0.64$^{+0.09}_{-0.12}$</td>
</tr>
<tr>
<td>Abell 2390</td>
<td>0.79$^{+0.38}_{-0.29}$</td>
<td>3.25$^{+1.77}_{-1.51}$</td>
<td>1250$^{+100}_{-275}$</td>
<td>0.76$^{+0.19}_{-0.12}$</td>
</tr>
<tr>
<td>Abell 1835</td>
<td>0.64$^{+0.21}_{-0.12}$</td>
<td>4.09$^{+0.54}_{-0.64}$</td>
<td>1275$^{+100}_{-150}$</td>
<td>0.50$^{+0.18}_{-0.09}$</td>
</tr>
<tr>
<td>MS2137-2353</td>
<td>0.18$^{+0.05}_{-0.02}$</td>
<td>8.30$^{+0.69}_{-1.22}$</td>
<td>800$^{+30}_{-30}$</td>
<td>0.16$^{+0.03}_{-0.02}$</td>
</tr>
<tr>
<td>RXJ1347-1145</td>
<td>0.40$^{+0.24}_{-0.12}$</td>
<td>5.87$^{+1.35}_{-1.44}$</td>
<td>1450$^{+300}_{-200}$</td>
<td>0.37$^{+0.12}_{-0.10}$</td>
</tr>
<tr>
<td>3C295</td>
<td>0.19$^{+0.07}_{-0.05}$</td>
<td>6.93$^{+1.67}_{-1.37}$</td>
<td>820$^{+90}_{-80}$</td>
<td>0.15$^{+0.07}_{-0.04}$</td>
</tr>
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<td>820$^{+90}_{-80}$</td>
<td>0.15$^{+0.07}_{-0.04}$</td>
</tr>
</tbody>
</table>

Figure 1. The observed (projected) spectrally-determined temperature profiles in the clusters, scaled in units of $kT_{2500}$ and $r_{2500}$, for the ACMD cosmology. PKS0745-191: open circles, Abell 2390: dark filled triangles, Abell 1835: grey filled stars, MS2137-2353: grey filled squares, RXJ1347-1145: dark filled circles, 3C295: open squares. The best-fit to the combined data set using the functional form in Section 3.1 is shown as the thin solid line.

Figure 2. The mass-temperature relation ($\Delta \chi^2 = 1.0$) for the NFW mass models. The best-fit values and 68 per cent confidence limits for $r_{2500}$, $M_{2500}$, $kT_{2500}$ and $L_{2500}$ are summarized in Table 3.

3 RESULTS

3.1 A universal temperature profile for relaxed clusters.

Fig. 1 shows the observed (projected), spectrally-determined temperature profiles in the clusters, in units of the gas mass-weighted temperature, $kT_{2500}$ and with the radial axis scaled in units of $r_{2500}$. The ACMD cosmology is assumed. The clusters exhibit similar scaled-temperature profiles which rise within $r \sim 0.3r_{2500}$ and then remain approximately isothermal up to $r_{2500}$. The combined data set can be modelled using a simple function of the form

$$T(r)/T_{2500} = T_0 + T_1(x/x_c)^\alpha/(1 + (x/x_c)^\alpha)$$

where $x = r/r_{2500}$, $T_0 = 0.40 \pm 0.02$, $T_1 = 0.61 \pm 0.07$, $x_c = 0.087 \pm 0.011$ and $\eta = 1.9 \pm 0.4$.

3.2 The mass-temperature relation

Fig. 2 shows the $M_{2500} - kT_{2500}$ relations for the SCDM and ACMD cosmologies. Fitting only the data for those clusters for which independent confirmation of the X-ray mass results is available from lensing studies (i.e. excluding 3C295)
Table 3. The total masses ($M_{2500}$, in units of $10^{14} M_\odot$), mean gas mass-weighted temperatures ($kT_{2500}$, in keV) and bolometric luminosities ($L_{2500}$, in units of $10^{45}$ erg s$^{-1}$) within radii $r_{2500}$ (in Mpc).

<table>
<thead>
<tr>
<th>Cluster</th>
<th>$E(z)$</th>
<th>$r_{2500}$</th>
<th>$M_{2500}$</th>
<th>$kT_{2500}$</th>
<th>$L_{2500}$</th>
<th>$E(z)$</th>
<th>$r_{2500}$</th>
<th>$M_{2500}$</th>
<th>$kT_{2500}$</th>
<th>$L_{2500}$</th>
</tr>
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<tr>
<td>PKS0745-191</td>
<td>1.158</td>
<td>0.89±0.04</td>
<td>6.06±0.10</td>
<td>9.50±0.16</td>
<td>7.35±0.23</td>
<td>1.050</td>
<td>0.68±0.03</td>
<td>4.96±0.58</td>
<td>9.55±1.06</td>
<td>4.26±0.11</td>
</tr>
<tr>
<td>Abell 2290</td>
<td>1.364</td>
<td>0.69±0.14</td>
<td>4.41±0.32</td>
<td>11.02±0.62</td>
<td>6.13±0.34</td>
<td>1.122</td>
<td>0.64±0.04</td>
<td>4.72±0.47</td>
<td>11.65±1.72</td>
<td>4.26±0.59</td>
</tr>
<tr>
<td>Abell 2385</td>
<td>1.401</td>
<td>0.72±0.03</td>
<td>5.41±0.14</td>
<td>11.05±1.91</td>
<td>8.50±0.19</td>
<td>1.135</td>
<td>0.66±0.02</td>
<td>5.23±0.45</td>
<td>11.23±1.03</td>
<td>5.61±0.03</td>
</tr>
<tr>
<td>MS2137-2353</td>
<td>1.505</td>
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<td>1.95±0.15</td>
<td>5.53±0.41</td>
<td>3.33±0.04</td>
<td>1.174</td>
<td>0.46±0.03</td>
<td>1.89±0.31</td>
<td>5.56±0.04</td>
<td>2.27±0.05</td>
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<tr>
<td>RXJ1347-1145</td>
<td>1.748</td>
<td>0.79±0.10</td>
<td>8.27±2.32</td>
<td>16.05±5.30</td>
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<td>8.95±3.37</td>
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<tr>
<td>3C295</td>
<td>1.765</td>
<td>0.64±0.03</td>
<td>1.63±0.34</td>
<td>5.51±0.67</td>
<td>1.74±0.08</td>
<td>1.279</td>
<td>0.41±0.04</td>
<td>1.60±0.33</td>
<td>5.61±0.75</td>
<td>1.30±0.06</td>
</tr>
</tbody>
</table>

Fig. 2. (a) The observed mass-temperature relation for the SCDM cosmology with $M_{2500}$ in $M_\odot$ and $kT_{2500}$ in keV. The solid line is the best-fitting power-law model $E(z)M_{2500} = A(kT_{2500}/10^6)$ with $A = 6.99 ± 0.57 \times 10^{14}$ and $\alpha = 1.5$ (fixed). The dotted curve is the predicted result from the hydrodynamical simulations of Evrard et al. (1996). (b) The results for the ACMD cosmology. The solid line is the best-fitting power-law model with $A = 5.38 ± 0.74 \times 10^{14}$ and $\alpha = 1.52 ± 0.36$. The dotted curve is the predicted result from the hydrodynamical simulations of Mathiesen & Evrard (2001).

ACDM cosmology. We conclude that the measured slope is consistent with the expected value of $\alpha = 1.5$ in all cases.

The dotted curves in Figs. 2a,b show the predicted (zero-redshift) relations for the SCDM and ACMD ($E(z)M_{2500} = 9.9 ± 1.5 \times 10^{14}kT^{1.5}$) cosmologies from the hydrodynamical simulations of Evrard et al. (1996) and Mathiesen & Evrard (2001), respectively. We have scaled the predicted curves from $\Delta = 500$ to $\Delta = 2500$ assuming $M_{2500} = M_{2500}(2500/500)^{-0.5}(T_{2500}/T_{500})^{1.5}$, which is consistent with the range of best-fit NFW mass models, and the SCDM simulations of Evrard et al. (1996; see their Table 5). Note that allowing for the presence of temperature gradients ($T_{2500} \neq T_{500}$) in the simulated clusters when applying this scaling does not affect the best-fit parameters for the theoretical $M_{2500} - kT_{2500}$ curves, since the curves are simply mapped onto themselves. For both SCDM and ACMD, the predicted normalization lies approximately 40 per cent above the observed value.

3.3 The temperature-luminosity relation

Fig. 3 shows the $kT_{2500} - L_{2500}$ relation for the ACMD cosmology. Fitting the $kT_{2500} - L_{2500}$ data for all six clusters using a power-law model of the form

$$\left(\frac{kT_{2500}}{10 \text{ keV}}\right) = B \left(\frac{L_{2500}}{10^{45} \text{ erg s}^{-1} \text{ E}(z)}\right)^{\beta}$$

and a $\chi^2$ estimator, we obtain $B = 0.43 ± 0.05$, $\beta = 0.45 ± 0.09$ for SCDM, and $B = 0.42 ± 0.05$, $\beta = 0.56 ± 0.10$ for ACMD. Using the BCES($X_2|X_1$) estimator of Akritas & Bershady (1996), which accounts for errors in both axes and the presence of possible intrinsic scatter, we obtain $B = 0.48 ± 0.07$, $\beta = 0.46 ± 0.08$ for SCDM and $B = 0.51 ± 0.05$, $\beta = 0.48 ± 0.06$ for ACMD.
We conclude that the slope of the temperature-luminosity relation for the present sample of hot, relaxed clusters is consistent with the predicted value of $\beta = 0.5$ (Section 1). Fixing $\beta = 0.33$ results in a poor fit: $\chi^2 = 12.5$ for 5 degrees of freedom, as opposed to $\chi^2 = 6.7$ with $\beta = 0.5$ ($\Lambda$CDM).

4 DISCUSSION

We have shown that within radii $r_{2500}$, corresponding to a fixed density contrast $\Delta = 2500$ with respect to the critical density at the redshifts of the clusters, the temperature profiles for the present sample of luminous, relatively relaxed lensing clusters exhibit an approximately universal form which rises within $r \sim 0.3 r_{2500}$ and then remains approximately constant out to $r_{2500}$. The enclosed masses, bolometric luminosities and mean mass-weighted temperatures within these radii scale in manner consistent with the predictions from the simple virial relations outlined in Section 1. We have confirmed the presence of a systematic offset of $\sim 40$ per cent between the normalizations of the observed and predicted $M_{2500} - kT_{2500}$ curves, in the sense that the predicted temperatures are too low for a given mass, for both the SCDM and $\Lambda$CDM cosmologies.

An important aspect of the present study is that independent confirmation of the X-ray mass measurements is available from gravitational lensing studies. For both Abell 2390 and RXJ1347-1145, the X-ray and weak lensing mass profiles are consistent within their 68 per cent confidence limits. For Abell 1835, 2390, MS2137-2353 and PKS0745-2390 and RXJ1347-1145, the X-ray and weak lensing mass limits. For Abell 1835, 2390, MS2137-2353 and PKS0745-2390, the X-ray and weak lensing masses are consistent within their 68 per cent confidence contours, although redshift measurements for the arcs (which are required to define the lensing masses precisely) are not available in all cases.

Thus, the presence of significant non-thermal pressure support (e.g. arising from turbulent and/or bulk motions and/or magnetic fields) can be excluded. We conclude that the systematic uncertainties associated with the individual mass measurements are small ($< 20$ per cent).

The offset between the observed and simulated mass-temperature curves cannot be explained by invoking an earlier formation redshift for the observed clusters (we assume that the clusters form at the redshifts they are observed) since, for the measured NFW mass distributions, $M_{2500}(z)$ drops as fast or faster than $E(z)$ rises as the formation redshift is increased. Our results suggest that on the spatial scales studied here, important physics may be missing from the reference simulations. One possible candidate is radiative cooling of the X-ray gas, which the Chandra data show to be significant within $r \sim 0.2 r_{2500}$ (e.g. Allen et al. 2001a,b,c; David et al. 2001; Schmidt et al. 2001), Pearce et al. (2000) show that the introduction of radiative cooling into their hydrodynamical simulations can lead to central temperature drops similar to those in Fig. 1. These authors also argue that cooling can lead to a significant increase in the mass-weighted temperature within $r \sim r_{2500}$ (as cooled, low-entropy gas is deposited and warmer, high-entropy material flows inwards and is compressed), which may be sufficient to account for the discrepancy between the observed and simulated curves. Detailed simulations of the $M_{2500} - kT_{2500}$ relation for large a sample of massive clusters, including the effects of radiative cooling, are required to address this issue.

The results presented in this paper should provide a useful calibrator for future studies of the X-ray properties of galaxy clusters. In future work we will examine the constraints that the present data place on radial variations in the X-ray gas mass fraction in the clusters and, therefore, $\Sigma_{\text{in}}$. We will also explore the ability of different parameterized mass models to explain the observed X-ray gas temperature and density profiles.

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REFERENCES

Arnaud, K.A., 1996, in Astronomical Data Analysis Software
and Systems V, eds. Jacoby G. and Barnes J., ASP Conf.
Series volume 101, p17