Investigating the EGRET-radio galaxies link with \textit{INTEGRAL}: the case of 3EG J1621 + 8203 and NGC 6251*

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Abstract. The analysis of an \textit{INTEGRAL} AO2 observation of the error contours of the EGRET source 3EG J1621 + 8203 is presented. The only source found inside the error contours for energies between 20 and 30 keV at $5\sigma$ detection significance is the FR I radio galaxy NGC 6251. This supports the identification of NGC 6251 with 3EG J1621 + 8203. The observed flux is higher and softer than observed in the past, but consistent with a variable blazar-like spectral energy distribution.

Key words. gamma-ray: observations – X–rays: galaxies – Galaxies: active – Galaxies: individual: NGC6251

1. Introduction

One of the long standing problems of modern high–energy astrophysics is to understand the nature of sources emitting the highest energy $\gamma$–rays. The Energetic Gamma–Ray Experiment Telescope (EGRET) Catalog, obtained from high–energy $\gamma$–ray observations ($E > 100$ MeV) performed between 1991 and 1995 (Hartman et al. 1999), has been revisited using deeper radio interferometric surveys at 8 GHz and improved radio spectral measurements in the northern and southern sky, down to $\delta > -40^\circ$ (Sowards-Emmerd et al. 2003, 2004). These interferometric data have allowed us to discriminate more efficiently between sources with a flat spectrum core from those with softer lobe emission. Combined with the associations proposed by Mattox et al. (2001) using older surveys at $\delta < -40^\circ$, a total of 128 AGN counterparts have been proposed. 60% of the identifications are considered very likely. Five associations with nearby Fanaro-Riley type I and II radiogalaxies (FRI and FRII, respectively), namely 3EG J0459 + 3352 (Sowards–Emmerd et al. 2003), 3EG J1324 − 4313 (Cen A, Steinle et al. 1998), 3EG J1646 − 0704 (Sowards–Emmerd et al. 2004), 3EG J1736 − 2908 (Di Cocco et al. 2004), 3EG J1735 − 1500 (Combi et al. 2003) and 3EG J1621 + 8203 (Mukherjee et al. 2002) open the exciting possibility of studying misaligned blazars in gamma rays. While blazars can be powerful $\gamma$–ray emitters because of relativistic beaming and amplification, the
analyzed data from X-ray satellites (ROSAT), the observer’s view or close to it. On the other hand, FR II radio galaxies are usually more intrinsically powerful than BL Lac objects and BL Lac objects can be seen either at the off-axis position, the source was not visible in the 20–30 keV energy range, but some explanations have been proposed (see the discussions in Sreekumar et al. 1999, and Mukherjee et al. 2002). The spectrometer SPI (Vedrenne et al. 2003) was performing an annealing and therefore no data are available. In addition, given the off-axis position, the source was not visible in the field of views of the X-ray monitor JEM-X (Lund et al. 2003) and the optical monitor OMC (Mas-Hesse et al. 2003), both also aboard INTEGRAL.

The analysis was performed on all the available data for an effective exposure toward 3EG J1621 + 8203 of 424 ks. The only known source detected (signal-to-noise ratio $S/N = 5\sigma$) inside the EGRET error contours is NGC 6251, found at coordinates ($J2000$) $\alpha = 16^h 33^m 10^s$ and $\delta = +82^\circ 35' 46''$ with an uncertainty of $S'$ at the 90% confidence level (Fig. 1). The count rate in the 20–30 keV energy band is $0.11 \pm 0.02 (1\sigma$ error) corresponding to a flux of $(6.7 \pm 2.1) \times 10^{-12}$ erg cm$^{-2}$ s$^{-1}$ (90% confidence level). Therefore, it is very interesting to search for other radio galaxies that may be counterparts of EGRET sources. Their identification with EGRET sources would have a large impact yielding, for instance, valuable constraints on the collimation of $\gamma$-ray emission from the jet and the energy of the jet particles. In addition, it would have a significant impact on population studies since the density of radio galaxies is far higher than that of blazars, and a significant reassessment of the origin of the diffuse extragalactic $\gamma$-ray background would be required.

Among the different cases available, we focus our attention on 3EG J1621 + 8203, that Mukherjee et al. (2002) have associated with the FR I radio galaxy NGC 6251 ($z = 0.02488$). They analyzed data from X-ray satellites (ROSAT, ASCA) and by cross–correlating with the NRAO VLA Sky Survey (NVSS), they found that, among the several sources found in the EGRET error contours, NGC 6251 is the most promising candidate to be a high–energy $\gamma$-ray emitter. However, it should be noted that inside the EGRET error contour there are several other X-ray and radio sources, some of them unidentified. Therefore, NGC 6251 was considered the best candidate among the known sources.

In order to resolve this doubt, we requested for a long observation with the IBIS telescope (Ubertini et al. 2003) on board the INTEGRAL satellite (Winkler et al. 2003). IBIS has a large field of view of $19^\circ \times 19^\circ$ at half response, with a high angular resolution of 12' sampled in 5' pixels in the low energy ($0.151 \text{–} 1\text{MeV}$) detector ISGRI (Lebrun et al. 2003) and in 10' pixels in the high energy ($0.175 \text{–} 10\text{MeV}$) layer PICsIT (Di Cocco et al. 2003).

The detection by EGRET of copious $\gamma$-rays from blazar jets implies that FR I and FR II radio galaxies are also intrinsically powerful $\gamma$-ray emitters, but their observation at large aspect angle should yield a much fainter flux and softer spectrum. This is indeed the case for Cen A (Steinle et al. 1998), that has a photon spectral index in the EGRET energy range ($E > 100\text{MeV}$) of $\Gamma = 2.40 \pm 0.28$, that appears to be slightly softer than the average value for blazars ($\Gamma = 2.15 \pm 0.04$), although still consistent within the measurement errors. Cen A has a large inclination angle ($\sim 70^\circ$), so that it is difficult to explain the detection of emission in the EGRET energy range, but some explanations have been proposed (see the discussions in Sreekumar et al. 1999, and Mukherjee et al. 2002).

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Here we report the analysis of the region containing the 3EG J1621 + 8203 error contours and the hard X-ray emission from NGC 6251.

Throughout the paper we adopted $H_0 = 70 \text{km s}^{-1} \text{Mpc}^{-1}$.

### 2. INTEGRAL data analysis

INTEGRAL observed 3EG J1621 + 8203 from June 17, 2004, 18:28 UT to June 23, 2004, 08:13 UT with a 5 x 5 dither pattern. The region of interest was about 4.6' off axis, because the observation of the EGRET source (Prop ID 0220018, PI Foschini) was amalgamated with the observation of the galaxy cluster Abell 2256 (Prop ID 0220020, PI Fusco Femia), which was the on-axis target.

The INTEGRAL data analysis described in the present work was done with the most recent version of the Offline Standard Analysis (OSA 4.1$^1$) whose algorithms for IBIS are described in Goldwurm et al. (2003).

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$^1$ Available through the INTEGRAL Science Data Centre (Courvoisier et al. 2003) at http://isdc.unige.ch/index.cgi?Soft+download

$^2$ The conversion of count rates in physical units has been done by normalizing to the count rates of the Crab Nebula, obtained in sev-
Table 1. Fluxes in the 20 – 30 keV energy band obtained in different epochs. In the case of ASCA and XMM-Newton, the value is extrapolated from the best fit model (see the notes for details). Columns: (1) Satellite name; (2) Date of the observation [DD–MM–YYYY]; (3) Photon index; (4) Flux [10^{-12} erg cm^{-2} s^{-1}]. The uncertainties in the parameters are at the 90% confidence level.

<table>
<thead>
<tr>
<th>Detector</th>
<th>Date</th>
<th>$\Gamma$</th>
<th>$F$ [10^{-12} erg cm^{-2} s^{-1}]</th>
</tr>
</thead>
<tbody>
<tr>
<td>ASCA</td>
<td>28 – 10 – 1994</td>
<td>1.8 ± 0.2</td>
<td>0.50 ± 0.06</td>
</tr>
<tr>
<td>BeppoSAX</td>
<td>19 – 07 – 2001</td>
<td>1.79 ± 0.06</td>
<td>1.73 ± 0.07</td>
</tr>
<tr>
<td>XMM-Newton</td>
<td>26 – 03 – 2002</td>
<td>1.91^{+0.08}_{-0.05}</td>
<td>1.19^{+0.08}_{-0.06}</td>
</tr>
<tr>
<td>INTEGRAL</td>
<td>17 – 06 – 2004</td>
<td>2.1^{+e}_{-0.1}</td>
<td>6.7 ± 2.1</td>
</tr>
</tbody>
</table>

a Sambruna et al. (1999). Flux extrapolated from the fit in the 0.6 – 10 keV energy band.
b Guainazzi et al. (2003), Chiaberge et al. (2003). Flux extracted from the joint fit LECS, MECS, PDS, in the 0.1 – 200 keV energy band.
c Gliozzi et al. (2004). Flux extrapolated from the fit in the 0.4 – 10 keV energy band.
d Present work.
e The conversion from rate to flux was done by assuming a Crab-like spectrum.

Confidence level). There is no detection in higher energy bands, with an upper limit (3$\sigma$) of 0.05 counts s$^{-1}$ in the 30 – 40 keV energy band, corresponding to a flux of 4 $\times$ 10$^{-12}$ erg cm$^{-2}$ s$^{-1}$ assuming a Crab-like spectrum.

Although the present statistics prevent a detailed spectral analysis, we performed some simulations with xspec (v. 11.3.1) and the IBIS/ISGRI RMF/ARF matrices available in the OSA 4.1 to understand the flux differences between INTEGRAL and BeppoSAX observations. The flux and photon index values observed by BeppoSAX corresponds to a count rate of about 0.03 c/s in the IBIS/ISGRI detector (20 – 30 keV energy range), to be compared with the value of 0.11 ± 0.02 from the present observation. This confirms that the flux variations are real.

Inside the EGRET probability contour there is only one other excess (marked with “A” in Fig. 1) with $S/N \approx 3.5\sigma$, but it does not correspond to any known source or even the unidentified X-ray sources found by Mukherjee et al. (2002). The closest source is 1RXS J163121.2+812641, 8' distant from the excess A. The low $S/N$ is not sufficient to claim a significant source detection, but it is rather likely to be a residual artifact of the deconvolution process.

The other detector layer of IBIS, PICsIT, had an overall exposure of 302 ks, but no detection was found, with an upper limit (3$\sigma$) of 1.2 $\times$ 10$^{-9}$ erg cm$^{-2}$ s$^{-1}$, in the 672 – 1036 keV band in the 3EG J1621 + 8203 region.

Fig. 2. Spectral energy distribution of NGC 6251 from Chiaberge et al. (2003) with the addition of the INTEGRAL/ISGRI data (horizontal bars are enlarged for clarity). The dashed line indicates the model fit to BeppoSAX data; the continuous line represents the model fit to INTEGRAL data. The new model parameters are as follows: source size $R = 3 \times 10^{16}$ cm, injected electron luminosity $L_{\text{ej}} = 6.7 \times 10^{44}$ erg/s, slope of the injection $p = 1.9$, $\gamma_{\text{min}} = 10^3$, $\gamma_{\text{max}} = 4 \times 10^5$, magnetic field $B = 0.04$ G, escaping time $t_{\text{esc}} = 15R/c$, beaming factor $\delta = 3.8$. See Chiaberge et al. (2003) for more details on the SSC model parameters.

3. Discussion

The identification of NGC 6251 as a counterpart of 3EG J1621 + 8203 is a complex puzzle, and with the present work we add some new pieces. To date Mukherjee et al. (2002) studied the ROSAT (0.1 – 2.4 keV) and ASCA (0.6 – 10 keV) observations, and the radio survey NVSS at 20 cm. They also investigated how the $\gamma$–ray emission from a relativistic jet could decrease as a function of the observer’s viewing angle and found that by assuming an angle of $\theta = 45'$ for the jet axis of NGC 6251, the flux emitted in the EGRET energy range is still detectable.

A complement, at radio wavelengths, to the work of Mukherjee et al. (2002) was recently published by Sowards–Emmerd et al. (2003), with a survey at 8.4 GHz, and they found that the only reliable counterpart of 3EG J1621 + 8203 is NGC 6251.

However, in addition to the spatial coincidence, it is necessary that the source is physically able to generate very high–energy $\gamma$–rays. In this respect, another piece in the puzzle was given by Guainazzi et al (2003) and Chiaberge et al. (2003): they found that the spectral energy distribution (SED) of NGC 6251 has two peaks, very similar to the typical SED of blazars (Fossati et al. 1998). Chiaberge et al. (2003) found also an upper limit to the jet direction of $\theta < 18'$ from the fit of the
SED with the synchrotron self–Compton (SSC) model (Fig. 4). On the other hand, Jones & Wehrle (2002) found $\theta < 47^\circ$ with the Very Long Baseline Interferometer (VLBI) measurement of the jet–counterjet ratio. Moreover, Chiaberge et al. (2003) evaluated also the beaming factor $\delta = 3.2$, that is lower than for blazars ($8 \leq \delta \leq 23$), but greater than the case of Cen A ($\delta = 1.2$).

The SSC model establishes a link between the radio and the hard X-ray emission, but before the present work, the only observation at energies greater than 10 keV was performed by the BeppoSAX PDS. However, there could be some doubt about the flux evaluation with the PDS, since it is a collimated detector with a field of view (FOV) of $1.3^\circ$ and there is an active Seyfert nucleus at a distance of $27^\prime$ from NGC 6251 (see the source n. 17 in the Fig. 3 of Mukherjee et al. 2002). Thanks to the unprecedented angular resolution of IBIS/ISGRI, it is possible to distinguish between the different contributions. The Seyfert nucleus is not detected in the present INTEGRAL observation and we can exclude any contamination. Moreover, IBIS/ISGRI shows that no other source or excess is present in the error contours of EGRET down to a flux of $\approx 5.4 \times 10^{-12}$ erg cm$^{-2}$ s$^{-1}$ ($4\sigma$, 20 – 30 keV energy range).

The SED of NGC 6251, updated with the data from the present observation, is shown in Fig. 4. The model SED has been derived using the numerical code described in Chiaberge & Ghisellini (1999). The data at different wavelengths are not obtained simultaneously, therefore, it is not possible to draw firm conclusions. However, with the INTEGRAL data it is possible to fit the data with the SSC, showing the source in a state with an overall higher flux. The parameters of the fit are similar to those in Chiaberge et al. (2003): in the present case the size of the source is slightly smaller by about a factor of less than 2 and the beaming factor increases from $\delta = 3.2$ to 3.8. Although the non-availability of simultaneous data prevents us from performing a detailed spectral modeling, we speculate that the “high state” observed during the present observation, is shown in Fig. 2. The model SED has been derived using the numerical code described in Chiaberge & Ghisellini (1999). The data at different wavelengths are not obtained simultaneously, therefore, it is not possible to draw firm conclusions. However, with the INTEGRAL data it is possible to fit the data with the SSC, showing the source in a state with an overall higher flux. The parameters of the fit are similar to those in Chiaberge et al. (2003): in the present case the size of the source is slightly smaller by about a factor of less than 2 and the beaming factor increases from $\delta = 3.2$ to 3.8. Although the non-availability of simultaneous data prevents us from performing a detailed spectral modeling, we speculate that the “high state” observed during the present observation, is shown in Fig. 2.

Another key question is still unexplained: the X-ray variability. It is indeed known that blazars display strong variability on different time scales, from hours to days (see e.g. Wagner & Witzel 1995, Ulrich et al. 1997). Table 4 shows the flux values obtained from recent observations with X-ray satellites. There are clear indication for both flux and photon index variabilities.

EGRET observed flux variations of 3EG J1621 + 8203 (see Nolan et al. 2003). The source is not persistent, i.e. it has been detected in individual EGRET observations, but not in the cumulated 4–yr data.

4. Conclusions

The findings of this INTEGRAL observation can be summarized as follows: NGC 6251 is the only hard X-ray (20 – 30 keV) source inside the EGRET error contours, down to a flux of $5.4 \times 10^{-12}$ erg cm$^{-2}$ s$^{-1}$ ($4\sigma$). The flux of the present observation is higher than was measured by BeppoSAX (and than the extrapolations into the hard X-ray band of the ASCA and XMM-Newton spectra), thus confirming that NGC 6251 is variable. The spectral energy distribution can be still modeled by the SSC model, although with slightly different parameters, but in agreement with a blazar-like behaviour.

These are two more pieces added to the puzzle of the identification of the EGRET source 3EG J1621 + 8203, but the final piece will be provided by the GLAST satellite, NASA’s next very high-energy $\gamma$–ray mission scheduled for launch in 2007, whose characteristics will assure a firm decision on the association of the high-energy $\gamma$–ray source 3EG J1621 + 8203 with NGC 6251.

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