A near-infrared and optical photometric study of the Sculptor dwarf spheroidal galaxy: implications for the metallicity spread

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ABSTRACT
The Sculptor dwarf spheroidal galaxy has a giant branch with a significant spread in colour, symptomatic of an intrinsic age/metallicity spread. We present here a detailed study of the Sculptor giant branch and horizontal branch morphology, combining new near-infrared photometry from the Cambridge InfraRed Survey Instrument (CIRSI), with optical data from the ESO Wide Field Imager. For a Sculptor-like old and generally metal-poor system, the position of Red Giant Branch (RGB) and Asymptotic Giant Branch (AGB) stars on the colour-magnitude diagram (CMD) is mainly metallicity dependent. The advantage of using optical-near infrared colours is that the position of the RGB locus is much more sensitive to metallicity than with optical colours alone. In contrast the horizontal branch (HB) morphology is strongly dependent on both metallicity and age. Therefore a detailed study of both the RGB in optical-near infrared colours and the HB can help break the age-metallicity degeneracy. Our measured photometric width of the Sculptor giant branch corresponds to a range in metallicity of 0.75 dex. We detect the RGB and AGB bumps in both the near-infrared and the optical luminosity functions, and derive from them a mean metallicity of $[\text{M}/\text{H}]= -1.3 \pm 0.1$. From isochrone fitting we derive a mean metallicity of $[\text{Fe}/\text{H}]= -1.42$ with a dispersion of 0.2 dex. These photometric estimators are for the first time consistent with individual metallicity measurements derived from spectroscopic observations. No spatial gradient is detected in the RGB morphology within a radius of 13 arcminutes, twice the core radius. On the other hand, a significant gradient is observed in the HB morphology index, confirming the ‘second parameter problem’ present in this galaxy. These observations are consistent with an early extended period of star formation continuing in time for a few Gyr.

Key words: galaxies: individual: Sculptor dwarf spheroidal – galaxies: stellar content – Local Group – infrared: stars

1 INTRODUCTION
The dwarf galaxies of the Local Group offer a unique opportunity to quantify the evolution and interactions of low-mass galaxies. Dwarf spheroidals were originally thought to be very similar in their metallicity and star formation histories to the galactic globular clusters, but their star formation history is now known to be more complex. The Sculptor dwarf spheroidal, the first dSph discovered (Shapley 1938), has a population which is predominantly old and moderately metal poor (e.g. Da Costa 1984, Mateo 1998, Monkiewicz et al. 1999, Dolphin 2002). However, the presence of an extended horizontal branch (e.g. Da Costa 1984, Majewski et al. 1999) and the detection of associated neutral hydrogen (Carignan et al. 1998) suggest the possibility of a more complex star formation history. A metallicity spread has been inferred in Sculptor from the large spread of its red giant branch (e.g. Da Costa 1984, Schweitzer et al. 1995, Kaluzny et al. 1993, Majewski et al. 1999), and the period distribution of RR Lyrae (Kaluzny et al. 1993, Kovacs 2001). Metallicity gradients have been discovered in several dwarf spheroidals (Harbeck et al. 2001). Majewski et al. (1999) and Hurley-Keller et al. (1999) showed that the red horizontal branch (RHB) of Sculptor is more strongly concentrated towards the centre than is the bluer population. Majewski et al. (1999) suggest this and the detection of two bumps in the Sculptor red giant branch are direct evidence for a bimodality in Sculptor’s metallicity distribu-
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data. In section 4 we present the detection of the RGB and
AGB bumps. The metallicity distribution function is derived
in section 5. Section 6 is devoted to the study of the varia-
tion of RGB and HB morphologies with radius. We conclude
with a discussion of the main results.

2 THE DATA
Photometric studies provide a valuable tool to determine in-
ternal metallicity spreads in dSph galaxies. Although indi-
vidual spectroscopic measurements are more precise, as yet
the numbers of stars with direct spectra is small. Most such
photometric studies use optical photometry. The combina-
tion of optical and near-infrared data provides substantially
enhanced information, reducing the effect of photometric er-
ors, allowing colour-colour selection of sources, and pro-
viding in particular the V-K colour, which is a good indicator
of the stellar effective temperature.

We have obtained wide-area near-infrared J and K-band
photometry of the Sculptor dSph galaxy, complementing this
with optical V and I-band data from the ESO archive.

Figure 1 presents the Sculptor area observed, while ta-
ble summarises the observations.

2.1 The near-infrared data
Near-infrared observations in the J(1.25µm) and
Ks(2.15µm) bands were made with the Cambridge Infrared
Survey Instrument (CIRSI) on the Las Campanas 2.5m duPont
telescope, and optical data from the ESO 2.2m Wide Field
Imager archive, has allowed us to undertake a broader wave-
band study of the metallicity spread in Sculptor. Observa-
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sion of the InfraRed Data Reduction (IRDR) software
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mary of the full process is given here. The updated ver-
sion of IRDR with its documentation are available at
http://www.ast.cam.ac.uk/~optics/cirsi/software A fuller
description is provided by Babusiaux & Gilmore (2005), in
their presentation of a CIRSI study of the Galactic bulge
and bar.

First each image is corrected for non-linearity. The dark
current for the relevant exposure time is then subtracted.
The data are flatfield corrected using lamp-on domeflats sub-
tracted with lamp-off domeflats and normalized to the first
detector sensitivity. These flatfields are also used to detect
bad pixels and create weight maps, used during the dither
frame coaddition.

The sky is subtracted in two passes. A first pass sky im-
age is made by median combining the nearest loop-combined
frames of the dither set. After a first dither frame coaddi-
tion, object masks are produced using SExtractor source
extraction (Bertin & Arnouts 1996). A masked frame is cre-
ated from this source-detection list, with an enlarged area
around each detected source being used. This object-masked
frame is used to make a second pass sky subtraction on each
loop.

Spatial offsets between loop-combined dither frames are
computed by cross-correlating object pixels mapped by SEx-

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2.2 The optical data

The optical data were obtained from the ESO 2.2m telescope. They consist of 3x300s dithered exposures in V and I. Each of the eight 4kx2k CCDs in the ESO 2.2m WFI covers around 8 arcmin x 16 arcmin of sky at a sampling of 0.238 arcsec per pixel, comparable with the CIRSI data. The total field of view of the WFI is about 33 arcmin x 33 arcmin, including small gaps of around 10 arcsec between CCDs. This field of view entirely covers the observed CIRSI fields.

The WFI data were processed using a variant of the standard optical pipeline described by Irwin 

2.1.2 CIRSI photometry

Once the data are reduced, we use the IRAF photometry routines. Source are detected using the IRAF DAOFIND procedure, with a significance threshold set at five-sigma. Aperture photometry is then obtained using the IRAF PHOT task.

The aperture radius is assigned for each dither frame to the measured PSF FWHM. Observations of standard stars from Persson et al. (1998) were obtained each night, and were used to derive the magnitude zero-point of each night. Standard star photometry used an aperture photometry radius of 20 pixels, equivalent to a diameter of 8 arcsec. The internal zero point dispersion derived from multiple observations of the standard stars during a night is 0.013 mag in J and 0.008 mag in K. The instrumental magnitudes derived using the psf-fwhm aperture are corrected for aperture effects using the curve-of-growth method (Stetson 1990), implemented with the IRAF MKAPFILE task applied to selected bright isolated stars. Airmass corrections of and are applied to the photometry. Finally, images detected near the borders of the images are eliminated, so that only stars observed in all the dither frames are kept. False detections located in the wings of highly saturated stars are manually deleted.

The photometric calibration was checked by correlating the brightest stars with the 2MASS catalogue. The CIRSI Ks photometry is consistent with the 2MASS photometric system. However, an offset in the J-band photometry of is observed. This is due to the 2MASS J-band filter being more extended into the atmospheric water absorption features at around 1.1 and 1.4 μm (Carpenter 2001) than the filter used for CIRSI.

The 5-σ magnitude completeness limits are about J~20 and Ks~18.8. The photometric errors are about 0.025 mag for J<17, 0.08 mag at J=20, 0.022 mag for Ks <16, and 0.08 mag at Ks =18.8.

## Table 1.

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<tr>
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<td>V</td>
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<td>3 x 300, dit-3</td>
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<tr>
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<td>1999-07-22</td>
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<td>0.91</td>
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Figure 2. The V-J vs J-Ks colour-colour diagram of all point sources detected in V, I, J and Ks. The solid line shows the quasar selection criterion from the KX method (Warren et al. 2000).

axis, was found to be a good approximation of the effect of the scattered light. Overall, this provides a zero-point calibration with 1-2% accuracy.

The 5-σ magnitude completeness limits are about V ∼24, I ∼22.5, so that the censorship on the data is due only to the CIRSI J and Ks photometric limits. The errors are smaller than 0.01 mag in V and 0.02 in I for magnitudes brighter than V ∼21.5 and I ∼19.5.

3 PHOTOMETRY OF THE SCULPTOR DSPH GALAXY

Our next task is to generate a list of point-sources which are stellar members of the Sculptor dSph galaxy. Extended objects are readily eliminated from further consideration using the morphological flags from the optical pipeline.

3.1 Selection of Sculptor stars

The VJK colour-colour diagram, figure 2, allows the detection of three other point-source populations unrelated to the Sculptor dSph galaxy. Red quasars are clearly seen with colours which are too red to correspond to any star. All sources with J-Ks > 1 were eliminated. A stream of stars with J-Ks about 0.8 mag and V-J colour redder than 3 mag can be seen. They are foreground low-mass dwarfs (e.g. Leggett 1992); these can be eliminated, without discarding any Sculptor stars, by excluding all stars redder than the colour of the tip of the Sculptor giant branch in all relevant colour-magnitude diagrams (e.g. V-I > 1.9). A number of probable blue quasars can be seen bluer in V-J than the main stellar locus of Sculptor and foreground stars. These are selected and eliminated according to the KX method criterion (Warren et al. 2000), which is shown as the solid line in figure 2. Remaining foreground Galactic stars are minimised by our selection inside each colour-magnitude diagram of the location of Sculptor member stars, as described further below.

Figure 3. Theoretical isochrones of the Padova group (a) AGB and RGB isochrones for an age of 14 Gyr and for metallicities [Fe/H] = -2.3,-1.7,-1.3,-0.7 dex from left to right. The data points are our photometry. (b) The RGB and main sequence turn-off for ages 14 and 8 Gyr and for the same four metallicities as in the top panel.

3.2 The Sculptor Colour-Magnitude diagrams

It is apparent from figure 3a, that there is a significant real width to the RGB of Sculptor in the (V-Ks,Ks) colour-magnitude diagram, confirming several previous studies of Sculptor CMD. The RGB photometric width is about ∆(V-Ks) = 0.3 at magnitude Ks =16, where the photometric measurement errors are 0.023 mag.

To determine the origin of this dispersion, theoretical isochrones from the Padova group, given in the ESO-WFI and 2MASS photometric systems by respectively Girardi et al. (2002) and Bonatto et al. (2004), have been overlaid on the Sculptor's CMDs, using a distance modulus (m-M)0 =19.54 and an extinction E(B-V)=0.02 (Mateo 1998). The extinction is derived in the different photometric bands using the Cardelli et al. (1989) extinction curve.

Figure 3b shows theoretical RGB isochrones for metallicities Z=0.0001, 0.0004, 0.001 and 0.004 with solar mixture, and ages 8 and 14 Gyr. It can be seen that the RGB is much more sensitive to metallicity than to age. The main sequence turn-off, which is more sensitive to age, was used by Monkiewicz et al (1999) to derive an age of 15±2 Gyr for Sculptor. The results of Dolphin (2002) and Rizzi et al. (2004) confirm the predominance of old stars in Sculptor. We therefore adopted isochrones of age 14 Gyr for figure 3b.
These isochrones clearly confirm the presence of a metallicity spread within the Sculptor RGB stars.

The colour magnitude diagrams can also be compared directly to globular cluster observations. Ferraro et al. (2000) and Saviane et al. (2001) provide fiducial lines for the RGB of Galactic globular clusters for a wide range of metallicity, in the (V,J,K) and (V,I) photometric systems. Figure 4 shows the fiducials of Ferraro et al. (2000) global clusters on the (V-Ks) colour-magnitude diagram of Sculptor. The transformation from absolute to relative magnitudes is the same as the one applied for the theoretical isochrones. However the V and K magnitudes are not on the same photometric system as ours, leading to expected photometric differences of the order of 0.1 mag. Considering that globular clusters tend to have alpha-enhanced element ratios, whereas the Sculptor stars do not, we indicate their global metallicities as defined by Ferraro et al. (1999). The definition of this global metallicity scale and how it can be translated into [Fe/H] for Sculptor is discussed in section 4. Here again a metallicity spread is confirmed as being consistent with the width of the Sculptor RGB. Figure 5 illustrates that all stars in Sculptor are more metal-poor than [M/H] = −1.0.

The spread of the RGB in the (J-Ks,Ks) and (V-I) CMDs is smaller and more sensitive to photometric errors than in (V-Ks,Ks). According to Saviane et al. (2000), a variation of metallicity from −2.0 to −1.5 dex results in a difference in V-I of 0.04 mag at 1=−2 (one magnitude brighter than the RGB bump), while according to Ferraro et al. (2000), the same variation of metallicity at Ks=−3 results in a variation of 0.2 mag in V-K. Considering the relative photometric errors in V, I and Ks, this means that V-Ks is 1.7 times as sensitive to metallicity as is V-I. We will therefore use preferentially the (V-Ks,Ks) CMD in the following to derive photometric metallicity indicators.

4 THE RGB AND AGB BUMPS

Local maxima are observed in the luminosity function of the giant branch of old metal-poor stellar populations. These RGB and AGB bumps are well known features of the colour-magnitude diagrams of globular clusters (e.g. Ferraro et al. 1999). Those two bumps are detected in all our visible and near-infrared bands (table 2), as illustrated in figure 4 for the V and Ks bands. To allow us to compare the absolute magnitudes of these features in Sculptor with other studies, the ESO WFI V and I photometry was converted into the standard Johnson photometry using the Girardi et al. (2002) isochrones in those two filter systems. At the location of the bumps, (V−I)V = 0.96, a given simulated star of the isochrones present the colours Vj − VWFJ = −0.06 and Ij − IFIJ = 0.1. As previously, the conversion to absolute magnitudes assumes a distance modulus of 19.54 mag and E(B-V)=0.02. Errors in the determination of the bump location are of about 0.1 mag.

Majewski et al. (1999) also detected those two bumps within the central 10′, but associated the second one with the RGB bump of a more metal-poor population. Indeed, the AGB bump of a population of metallicity [Fe/H]≃ −1.5 is located at the same position on the CMD as is the RGB bump of a population of metallicity [Fe/H]≃ −2. The V magnitude of the second bump is consistent with the value of MV(AGBbump) = −0.3 ± 0.1 used by Ferraro et al. (1999). Its clear detection can be explained by the fact that the luminosity level of the AGB bump stays fairly constant with the cluster metallicity (e.g. Castellani et al. 1994). The AGB bump being always bluer than the RGB, it explains the detection of this second bump on the blue side of the RGB by Majewski et al. (1999). Using, as they did, a division of
the giant branch into a red and blue part (figure 6). We not only find as expected the second bump on the blue side of the giant branch, but also the RGB bump shifted by about 0.2 magnitudes (figure 4b), which is consistent with a variation in metallicity. If the second bump is to be due to a distinct metal-poor population, being clearly detected at all wavelengths, this second population should also show a clear imprint on the tip of the red giant branch. No such distinct second RGB can be detected (figures 6a. and 6d). We then conclude that the second bump is the AGB bump.

The V magnitudes of the HB and the RGB bump are good indicators of the metallicity (e.g. Ferraro et al. 1999). The peculiar shape of Sculptor’s HB however makes the determination of its mean V magnitude unreliable (figure 7). The V magnitude of the RGB bump leads to a global metallicity of [M/H]=−1.30±0.12 from the calibration of Ferraro et al. (1999). Its Ks magnitude leads to [M/H]=−1.39±0.14 according to Cho & Lee (2002), while it leads to [M/H]=−1.19±0.12 from the new calibration of Valenti et al. (2004). Those calibrations were made using the relation of Salaris et al. (1993):

\[
[M/H] = [\text{Fe}/H] + \log(0.638 + 10^{[\alpha/\text{Fe}]}) + 0.362
\]  

(1)

Sculptor does not seem to present a strong enhancement of alpha-elements: Shetrone et al. (2003) measured for 5 stars the following values for [\alpha/\text{Fe}]: 0.18, 0.13, 0.13, −0.01 and 0.23 (table 2 of Tolstov et al. 2003). For [\alpha/\text{Fe}]=0.13, [M/H]=−1.3 can be translated to [\text{Fe}/H]=−1.4 by equation 1.

It should be stressed that the indicated errors do not take into account the uncertainty in the distance modulus, estimated to be 0.08 mag in Mateo (1998). Another metallicity indicator, independent of the distance modulus and of zero point calibration errors, is the difference between the AGB and RGB bump luminosities. From Ferraro et al. (1999), with ΔV_{RGB bump} = 0.52 ± 0.14, we can derive [M/H]=−1.4 ± 0.2.

Those metallicity indicators agree with the previous comparisons of the RGB morphology with theoretical isochrones and globular clusters (figure 6a. and figure 6d).

5 THE METALLICITY DISTRIBUTION FUNCTION

The mean fiducial of the Sculptor red giant branch was computed through a least squares fit to a second order polynomial on the (V−Ks,Ks) CMD. Horizontal branch stars and foreground stars have been eliminated by selecting only the stars 0.2 magnitudes away from the mean fiducial (figure 6). Only the giant branch stars brighter than Ks=18.7 and up to the RGB tip (Ks>13.8) have been selected. As the AGB stars occupy the same location on the CMD as the most metal poor RGB population, no attempt to discard AGB stars was made.

We derive the metallicity distribution of those selected stars from a function of the metallicity of the isochrones:

\[
\text{V}−K_s = a_0 + a_1 K_s + a_2 K_s^2
\]  

(2)

A second order polynomial regression of those coefficients as a function of the metallicity of the isochrones is obtained:

\[
a_i = a_{i(0)} + a_{i(1)} [\text{Fe}/H] + a_{i(2)} [\text{Fe}/H]^2
\]  

(3)

By inverting equation 2 each point in the (V−Ks,Ks) CMD can then be assigned an estimate of its metallicity. Taking into account only the uncertainty of the accuracy on the polynomial regression of equations 2 and 3 and the photometric errors, the typical uncertainty in the resulting measurement of the metallicity of a individual star is smaller than 0.04 dex.
A near-IR study of the Sculptor dwarf spheroidal

Figure 8. Photometric metallicity distribution function as derived using theoretical isochrones.

The resulting metallicity distribution function is represented in figure 8. The secondary peak at [Fe/H] = −2.2 is an artefact due to AGB stars and should be ignored. The mean metallicity obtained is [Fe/H] = −1.42 with a dispersion of 0.2 dex. This mean metallicity is in agreement with the value obtained from the RGB bump. A comparison of those metallicity estimates with the literature is given in the discussion section.

6 METALLICITY GRADIENT INDICATORS

6.1 The Giant Branch

Since the RGB in our photometric system provides a good indicator of metallicity, we studied its variation with radius as a test of a possible metallicity gradient in Sculptor.

The metallicity indicators derived in the previous section have been studied as a function of radius. The central 10 arcminutes of Sculptor has zero ellipticity (Irwin & Hatzidimitriou 1995), so as our data do not extend beyond a radius of 15 arcminutes, we use a simple circular annulus. Figure 9 shows that no metallicity gradient is detected within our data. This gives an upper limit of 0.03 dex for the metallicity gradient within twice the core radius of Sculptor.

6.2 The Horizontal Branch

The horizontal branch of Sculptor can be clearly divided into a blue (B) and a red (R) part, lying on either side of the instability strip (V). The ratio of the number counts of those different parts, quantified by the HB index (B-R)/(B+V+R), is dependent on the metallicity. But there is a well known ‘second parameter problem’, which could be age (e.g. Lee et al. 1994).

The HB index was computed from the (V-I,V) CMD, as illustrated in figure 7, for the same radial annuli as used for the RGB study. Both the full ESO WFI field of view data and the sub-area in common with the CIRSI field of view are presented in figure 10. A K-S test for the hypothesis that the red and blue horizontal branch stars have the same radial distribution gives a significance level of 10^{-9}%. This confirms the HB gradient detected by Hurley-Keller et al. (1999) and Majewski et al. (1999).

The Lee et al. (1994) models give theoretical isochrones that show, for an HB index between -0.5 and 0.5 and a given age, a linear relation between the HB index and the metallicity:

\[ [Fe/H] \approx -0.34 \times HB_{\text{index}} + \text{cte} \] (4)

The observed gradient in HBindex of about 0.5 then corresponds to a gradient in metallicity of 0.17 dex. Considering the upper limit of 0.04 dex for a metallicity gradient derived previously from the RGB morphology, an age gradient is required to explain the observed HB gradient. As always when discussing HB morphology however, one must recall that the ‘second parameter problem’ is not yet solved and that another parameter may influence the HB morphology.

The simplest conclusion is that a small gradient in mean age is apparent in Sculptor and that an eventual small metallicity gradient associated with it would have a too small effect on the RGB compared to the large abundance dispersion to be detected.
The combination of near-infrared photometry from CIRSI with optical data from the ESO WFI, allowed a detailed study of the Sculptor dwarf spheroidal giant branch morphology. We confirm that the broad giant branch of Sculptor demonstrates an intrinsic metallicity spread (e.g. Da Costa 1984, Schweitzer et al. 1995, Kaluzny et al. 1995, Majewski et al. 1999). From our photometric study we quantify this spread into a metallicity range of $\Delta [\text{Fe/H}] = 0.75$ dex. The RGB and AGB bumps are detected in all the optical and near-infrared luminosity functions, excluding the substantial metal-poor contribution to Sculptor’s metallicity distribution proposed by Majewski et al. (1999). We derive a mean metallicity within two core radii in Sculptor of about $[\text{Fe/H}] = -1.4$ from both the RGB and AGB bumps magnitudes and isochrones fitting. Our mean metallicity and the metallicity range are higher than those derived in Da Costa (1984) from photometry of the RGB and in Kaluzny et al. (1995) from RR Lyrae stars: these results were summarized in the Mateo (1998) review as a mean of $[\text{Fe/H}] = -1.8 \pm 0.1$ with a spread of 0.3 dex. However our photometry is in agreement with the metallicity estimations from Sculptor RR Lyrae stars of Kovacs (2001) and CaII triplet observations of 37 stars of Tolstoy et al. (2001). It is in excellent agreement with the very recent spectroscopic survey of Tolstoy et al. (2004), whose derived metallicity distribution inside two core radii indicates a mean metallicity of $-1.4$ dex, and metallicity range of about 1 dex. Their data show the population structure is complex, with the more metal-poor part of the distribution function becoming dominant at radii beyond those we have studied here.

We do not detect a gradient in the RGB morphology within a radius of 13′, 2.2 times the Sculptor core radius. Although Harbeck et al. (2001) and Tolstoy et al. (2004) find a metallicity gradient in Sculptor, our result is in agreement with their data. Indeed figure 6 of Harbeck et al. (2001) shows that the radial distribution of blue and red RGB stars begins to differ only after 13′. Figure 3 of Tolstoy et al. (2004), based on spectroscopic data, indicates that a metallicity gradient is indeed visible only beyond this radius. On the other hand, we do detect at high significance a gradient in the horizontal branch (HB) morphology, confirming the results of Hurley-Keller et al. (1999) and Majewski et al. (1999). As this cannot be explained by metallicity, the most likely second parameter could be age. Hurley-Keller et al. (1999) did not find evidence for a gradient in the main-sequence turn off, leading to an upper limit of a 2 Gyr variation at constant metallicity. According to the models of Lee et al. (1994) a small variation in age of even a few Gyr leads to a strong variation in the HB morphology. Age could then still be the second parameter. Moreover, [Fe/H] and age are not the only variables which can affect the HB morphology, but also other element abundances, in particular the $[\text{O/Fe}]$ ratio. Lee et al. (1994) and Hurley-Keller et al. (1999) detected another population presenting a gradient in Sculptor: a ‘spur’ of stars extending ~0.7 mag above the old main sequence turn off. They conclude that it cannot be explained by the presence of an intermediate age population, preferring the interpretation of the spur as a binary sequence, and speculate that it could be related to the variation of the HB morphology.

No significant intermediate age population has been found in Sculptor, excluding star formation within the last 5 Gyr. The large metallicity spread requires extended star formation, while the evidence of low alpha-element enhancement implies extended star formation and self-enrichment over a period of at least 1 Gyr: quite long enough to affect the HB morphology. The HB morphology gradient implies that the most recent star formation episode occurred in the centre of the galaxy, consistent with naive expectation that gas is more easily retained deeper in the galaxy potential well. Indeed in most dwarf galaxies observed with sufficiently deep wide-field imaging, populations gradients have been found with the younger stars being more centrally concentrated (e.g. Saviane et al. 2001 and references therein). Our lack of detection of a metallicity gradient may be explained by the age-metallicity degeneracy that would hide a small age and metallicity difference and by the stronger dependence of the HB on other parameters such as the oxygen abundance, stellar rotation or binarism, as already suggested by Hurley-Keller et al. (1999) and Majewski et al. (1999).

All this could be consistent for Sculptor with a single period of star formation extended in time for of the order of a few Gyr. Tolstoy et al. (2003) conclude that their study of the element abundances are consistent with a closed-box chemical evolution scenario. The small dynamical mass of dwarf spheroidals such as Sculptor means that their binding energy is small compared to the energy released by several supernovae, which leads the high metallicity spread and relatively high mean metallicity derived for Sculptor puzzling: how did the gas stay bound long enough to have an extended star formation and gas enrichment? The star formation rate should be low to allow the chemical enrichment to proceed gradually. Hydrodynamical simulations are trying to answer this question (e.g. Carraro et al. 2001, Carigi et al. 2002).

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REFERENCES


