FAR-INFRARED CHARACTERIZATION OF AN ULTRA-LUMINOUS STARBURST ASSOCIATED WITH A MASSIVELY-ACCRETING BLACK HOLE AT $z = 1.15$

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ABSTRACT

As part of the All Wavelength Extended Groth Strip International Survey (AEGIS), we describe the panchromatic characterization of an X-ray luminous active galactic nucleus (AGN) in a merging galaxy at $z = 1.15$. This object is detected at infrared (8 $\mu$m, 24 $\mu$m, 70 $\mu$m, 160 $\mu$m), submillimeter (850 $\mu$m) and radio wavelengths, from which we derive a bolometric luminosity $L_{\text{bol}} \sim 9 \times 10^{12} L_{\odot}$. We find that the AGN clearly dominates the hot dust emission below 40 $\mu$m but its total energetic power inferred from the hard X-rays is substantially less than the bolometric output of the system. About 50% of the infrared luminosity is indeed produced by a cold dust component that probably originates from enshrouded star formation in the host galaxy. In the context of a coeval growth of stellar bulges and massive black holes, this source might represent a “transition” object sharing properties with both quasars and luminous starbursts. Study of such composite galaxies will help address how the star formation and disc-accretion phenomena may have regulated each other at high redshift and how this coordination may have participated to the build-up of the relationship observed locally between the masses of black holes and stellar spheroids.

Subject headings: galaxies: high-redshift — infrared: galaxies — cosmology: observations

1. INTRODUCTION

Galaxies with a bolometric luminosity exceeding $10^{12} L_{\odot}$ are often powered by a combination of massive star formation and accretion of material around active nuclei (e.g., Genzel et al. 1998). Early in cosmic history the connection between these two phenomena may have led to a coeval growth of super massive black holes (SMBHs) and stellar spheroids (e.g., Page et al. 2001), and could thus be the foundation of the correlation observed between the masses of these two components in the local Universe (e.g., Gebhardt et al. 2000). However, the implication of such a co-evolution in the more general context of the stellar mass built-up history (e.g., Dickinson et al. 2003) and SMBH formation (e.g., Barager et al. 2003) has not been fully addressed. Understanding the importance of this “coordinated” activity of starbursts and active galactic nuclei (AGNs) requires properly deconvolving their respective contributions within individual objects, a challenging task particularly for distant sources.

Because of the complexity of the AGN and starburst spectral energy distributions (SEDs), this decomposition can best be achieved by combining data from wavebands that offer distinctive spectral features to characterize these phenomena. For instance, active nuclei are strong emitters at high energy and they are usually associated with a continuum of hot dust peaking in the mid-infrared (IR). Luminous starbursts, on the other hand, are characterized by a colder dust component and emit the bulk of their luminosity in the far-IR. To illustrate how the coincidence of these two processes at high redshift could be studied using larger samples, we present a panchromatic analysis of a luminous AGN referenced as CXO–GWS–J141741.9+522823 (hereafter CXO-J1417) by Nandra et al. (2004). It is embedded in an ultraluminous infrared galaxy (ULIRG) at $z = 1.15$ and its detection across the full electromagnetic spectrum allows us to constrain the level of star formation also present in this source. We assume a ΛCDM cosmology with $H_0 = 70$ km s$^{-1}$ Mpc$^{-1}$, $\Omega_m = 0.3$ and $\Omega_{\Lambda} = 0.7$.

2. THE DATA

CXO-J1417 is a bright X-ray source also known in the literature as CFRS 14.1157 or CUDSS 14.13. It is associated with a very red galaxy ($I-K_{AB} \sim 2.6$; Webb et al. 2003; see also Wilson et al. 2006) detected in the mid-IR (Flores et al. 1999; Higdon et al. 2003; Barlow et al. 2006; Ashby et al. 2006), submillimeter (Eales et al. 2000; Webb et al. 2003; see also Wilson et al. 2006) and radio (Fomalont et al. 1991; Chapman et al. 2003). High resolution images obtained with HST (Davis et al. 2006; Lotz et al. 2006) reveal several components presumably interacting with each other (see Fig. 1). The brightest is dominated

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by a point source located at RA = 14h17m41s58 and Dec = 52°28′23.65″ (J2000, δRA ∼ δDec ∼ 0.07″) coinciding precisely with the position of the X-ray detection (Mivazj et al. 2004).

At this location there is also a bright and point-like object detected at 70 μm and 160 μm. In spite of the large beam used for the data at long wavelengths, the fact that the mid-IR counterpart of CXO–J1417 is a 2 orders of magnitude brighter than any other galaxies detected at 24 μm in this area strongly suggests that this far-IR emission is also associated with the X-ray source. Its fluxes in the MIPS 24/70/160 μm bands were measured via PSF fitting. They are reported in Table 1, which also summarises the full multi-wavelength photometry of the object (see Davis et al. 2000 for a description of our data set). Optical and near-IR fluxes were measured within a 1″-radius aperture centered at the position of CXO–J1417. With the exception of the GALEX data where we believe the mid-IR counterpart of CXO–J1417 is two orders of magnitude brighter than any other galaxies detected at 24 μm in this area strongly suggests that this far-IR emission is also associated with the X-ray source.

A redshift of z = 1.15 was reported by Hammer et al. (2004) based on UV/optical spectroscopy. In Fig. 2 we display the combined spectrum obtained at Keck by Davis et al. (2003) using LRIS and DEIMOS. The presence of an AGN is clearly confirmed by the detection of [NeV] and MgII. There is however considerable selfabsorption of the latter (Sarajedini et al. 2003), which prevents firm classification as a type 1 or a type 2 object. Interestingly, we also note the detection of Ca K+H absorption lines redshifted by ~150-200 km s⁻¹ relative to [OII] and which suggests the presence of gas inflow in the galaxy.

Finally, mid-IR spectroscopy was carried out by Higdon et al. (2004) as well as our own group. We obtained 17 900 s exposure time for each of the two SL1: 7.4–14.5 μm; SL2: 5.2–8.7 μm) and each of the two Long Low modules of the IRS (LL1: 19.5–38 μm; LL2: 14.0–21.3 μm), see Higdon et al. (2004) for a description of our data set. Optical and near-IR fluxes were measured within a 1″-radius aperture centered at the position of CXO–J1417. With the exception of the GALEX data where we believe the mid-IR counterpart of CXO–J1417 is two orders of magnitude brighter than any other galaxies detected at 24 μm in this area strongly suggests that this far-IR emission is also associated with the X-ray source.

The full spectral energy distribution of CXO–J1417 is illustrated in Fig. 3. It has a very red continuum with a steep rise from the optical up to the mid-IR, and the rest-frame 2.5–16.5 μm bands were measured via PSF (solid line) reveals that the central core is not resolved by HST. The object lying 1.4″ to the South is a foreground galaxy at z = 1.00.

### TABLE 1

<table>
<thead>
<tr>
<th>Band</th>
<th>Flux/Flux density</th>
<th>Reference</th>
</tr>
</thead>
<tbody>
<tr>
<td>2–10 keV</td>
<td>3.8±0.3×10⁻¹⁷ erg cm⁻² s⁻¹</td>
<td>1</td>
</tr>
<tr>
<td>0.5–2 keV</td>
<td>1.3±0.1×10⁻¹⁷ erg cm⁻² s⁻¹</td>
<td>1</td>
</tr>
<tr>
<td>FUV (1539 Å)</td>
<td>&lt;0.45 μJy (3σ)</td>
<td></td>
</tr>
<tr>
<td>NUV (2316 Å)</td>
<td>0.5±0.1 μJy</td>
<td></td>
</tr>
<tr>
<td>B (4389 Å)</td>
<td>0.8±0.1 μJy</td>
<td></td>
</tr>
<tr>
<td>R (6601 Å)</td>
<td>6.0±0.1 μJy</td>
<td></td>
</tr>
<tr>
<td>I (8133 Å)</td>
<td>16.7±0.8 μJy</td>
<td></td>
</tr>
<tr>
<td>J (1.2 μm)</td>
<td>57.6±0.5 μJy</td>
<td></td>
</tr>
<tr>
<td>K (2.2 μm)</td>
<td>117±1 μJy</td>
<td></td>
</tr>
<tr>
<td>IRAC 3.6 μm</td>
<td>580±0.4 μJy</td>
<td>2, 3</td>
</tr>
<tr>
<td>IRAC 4.5 μm</td>
<td>981±0.5 μJy</td>
<td>2, 3</td>
</tr>
<tr>
<td>IRAC 5.8 μm</td>
<td>1448±4 μJy</td>
<td>2, 3</td>
</tr>
<tr>
<td>IRAC 8.0 μm</td>
<td>2225±4 μJy</td>
<td>2, 3</td>
</tr>
<tr>
<td>IRS 16 μm</td>
<td>33±0.7 mJy</td>
<td>4</td>
</tr>
<tr>
<td>MIPS 24 μm</td>
<td>5.75±0.1 mJy</td>
<td></td>
</tr>
<tr>
<td>MIPS 70 μm</td>
<td>20.1±1.2 mJy</td>
<td></td>
</tr>
<tr>
<td>MIPS 160 μm</td>
<td>105±30 mJy</td>
<td></td>
</tr>
<tr>
<td>SCUBA 850 μm</td>
<td>3.3±1 mJy</td>
<td>5</td>
</tr>
<tr>
<td>VLA 5 GHz</td>
<td>53±4 μJy</td>
<td>6</td>
</tr>
<tr>
<td>VLA 1.4 GHz</td>
<td>110±40 μJy</td>
<td>6, 7, 8, 9</td>
</tr>
</tbody>
</table>

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**Footnotes:**

16 The FWHM of the MIPS PSF is 18″ and 40″ at 70 μm and 160 μm respectively.

17 General Observer program ID 3216.
X-ray data point therefore to a luminous AGN characterized by substantial obscuration.

This active nucleus is probably not the only source powering the bolometric luminosity of CXO–J1417. Although they are strongly diluted by the continuum emission from the AGN, the mid-IR broad bands from the Polycyclic Aromatic Hydrocarbons (PAHs) often seen in star-forming environments seem to be detected in our IRS spectrum (see Fig. 2b). Assuming a power-law continuum superimposed with a typical PAH template and leaving the redshift of the latter as a free parameter, we generated a series of simulated spectra that we compared to our data. The $\chi^2$ shows a clear minimum when the PAH component is shifted to the distance of CXO–J1417, which suggests that these features are detected with relatively good confidence. Furthermore, the far-infrared and submillimeter detections reveal a very luminous cold dust component typical of those observed in dusty star-forming galaxies (see Sect. 4). To quantify its contribution relative to the much warmer dust seen in the mid-IR, we decomposed the global IR SED beyond 1 μm into (i) a single modified black body accounting for the cold component, characterized by a temperature $T \sim 40$ K and a dust emissivity, and (ii) a combination of several blackbodies with temperatures ranging from 150 K to 2000 K and reproducing the warm dust emission (see Fig. 3). We derived $1-1000 \mu m$ integrated luminosities of $4.5 \times 10^{12} L_\odot$ for each of these two components. This leads to a total IR luminosity $L_{1-1000\mu m} = 9.0 \pm 0.4 \times 10^{12} L_\odot$, where the uncertainty is driven by the determination of the temperature and the emissivity of the cold dust emission.

CXO–J1417 is also remarkably quiet in the radio. The spectral index ($\alpha \sim 0.6$) is typical of synchrotron emission from starbursts, though we cannot exclude a flatter continuum because of the substantial uncertainty on the observed flux density at 1.4 GHz. Assuming the standard far-IR/radio correlation, we would infer an IR luminosity $L_\nu > 2.5\times10^{12} L_\odot$ lower than that implied by our fit of the cold dust component. As it has already been observed at low redshift (e.g., Rieke et al. 1980; Clemens & Alexander 2004; Gallimore & Beswick 2004) this radio faintness could result from substantial free-free absorption in the interstellar medium of the galaxy. It could also be due to a synchrotron deficiency characteristic of a very recent episode of star formation (Roussel et al. 2003).
4.1. On the nature of CXO–J1417

Although the bolometric correction for luminous and strongly-absorbed X-ray sources has not been very well constrained so far, the obscuration toward CXO–J1417 is still reasonable enough to allow a fairly secure estimate of the total luminosity of the active nucleus in this object. Comparison of the SED shown in Fig. 3 with typical AGN templates (e.g., Elvis et al. 1994, Silva et al. 2004) indicate that this bolometric luminosity should typically range between $1 \times 10^{12} L_\odot$ and $3 \times 10^{12} L_\odot$. While the AGN can thus power most of the hot dust detected in the system, it is not energetic enough to account also for the far-IR emission. We argue that the cold component is produced by a deeply enshrouded starburst in the host galaxy (see also e.g., Waskett et al. 2003). CXO–J1417 could be therefore a high redshift analog of some nearby “composite” ULIRGs where the contribution of the AGN to the bolometric output is comparable to that of the star-forming activity (Farrah et al. 2003). It could also be similar to other distant X-ray selected sources that were detected at long wavelengths (e.g., Page et al. 2000), though this far-IR emission has been sometimes assumed to originate from the active nucleus rather than star formation (Barger et al. 2005).

Assuming the calibration from Kennicutt (1998), the IR luminosity of the cold component translates into a star-formation rate $SFR \sim 750 M_\odot$ yr$^{-1}$. Such enhanced levels of activity usually occur within embedded and very compact regions surrounding the cores of galaxies ($\sim 100–300$ pcs, Soifer et al. 2000). It is consistent with the absence of direct star formation signatures as inferred from our UV/optical photometry, as well as from the ACS image taking into account the spatial resolution of the HST data (i.e., $\sim 1$ kpc at $z = 1.15$).

4.2. Implications

Assuming a typical accretion efficiency $\epsilon = 0.1$ (Marconi et al. 2004), the luminosity of the AGN in CXO–J1417 translates into a mass accretion rate $dM/dt \sim 3.1 M_\odot$ yr$^{-1}$. This is typical of quasars at $z \sim 1$ (McLure & Dunlop 2004), but it is larger than the rates measured in sources experiencing similar levels of starburst activity such as the more distant SCUBA sources (Alexander et al. 2003). Furthermore, the bolometric luminosity of this object and the obscuration toward its nucleus suggest that the gas fueling and the accretion are occurring quite efficiently, probably close to the Eddington limit. Under this hypothesis we would derive a black hole mass of $\sim 1.4 \times 10^8 M_\odot$. This is typically an order of magnitude larger than the mass of the SMBHs determined in the submillimeter galaxies, and it would be even larger in the case of a sub-Eddington accretion. These properties suggest that CXO–J1417 is an object sharing characteristics with both starburst-dominated galaxies and quasars, where violent star formation is still happening while a massive black hole has already formed.

High redshift ULIRGs showing a mixture of star formation and AGN such as CXO–J1417 could be interesting as tests of the evolutionary sequences that have been proposed to understand the connection between the two phenomena (e.g., Sanders et al. 1988). In such scenarios for instance, merging galaxies first trigger powerful star formation, and as material settles into the cores of these objects it feeds a supermassive black hole that eventually emerges as a luminous AGN. The latter can then produce strong winds and outflows that feed energy back into the surrounding galaxy and may either quench or reactivate star formation (Springel et al. 2005, Hopkins et al. 2003). In the case of CXO–J1417 there is a dominant contribution of the nucleus in the near-IR and it is not clear whether an underlying bulge has already formed in the host galaxy. However, sources experiencing star formation and disk-accretion that both radiate a similar amount of energy throughout their lifetime would evolve toward massive galaxies that lie significantly out of the local $M_\text{BH} - \sigma$ relationship (Page et al. 2001).

Such transitional cases might be rare locally (e.g., Genzel et al. 1998). At higher redshift however, their importance relative to the infrared/submillimeter or X-ray selected objects where one type of activity (i.e., star formation or accretion) largely dominates is not yet known. Interestingly, CXO–J1417 lies at the knee of the 2–8 keV luminosity function derived by Barger et al. (2003) at $0.8 \leq z \leq 1.2$ but it is much more IR-luminous than most of star-forming galaxies at this epoch of cosmic history (Le Floc’h et al. 2005). Searching for similar objects at higher redshifts when ULIRGs were a major component of the starbursting population (Blain et al. 2002) should allow us to explore in more detail the role that this co-existence of AGNs and starbursts within galaxies played in shaping the present-day Universe. Even though their identification could be challenging, CXO–J1417 points to the type of evidence required for this goal. Large data sets from existing surveys like AEGIS should provide this information for enough sources to probe the prevalence of this phase of galaxy evolution at $z \geq 1$.

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REFERENCES


On the coeval growth of bulges and massive black holes